

**What can we infer
from cosmology
for fundamental laws
of physics ?**

Three lessons

- 1) Fundamental “constants” are not constant
- 2) No conflict between quantum physics and general relativity
- 3) Time and space are born at the big bang

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**Fundamental
“constants” are not
constant**

*Have coupling constants in the
early Universe
other values than today ?*

Yes !

Fundamental couplings in quantum field theory

*Masses and coupling constants
are determined by properties
of **vacuum** !*

Similar to Maxwell – equations in matter

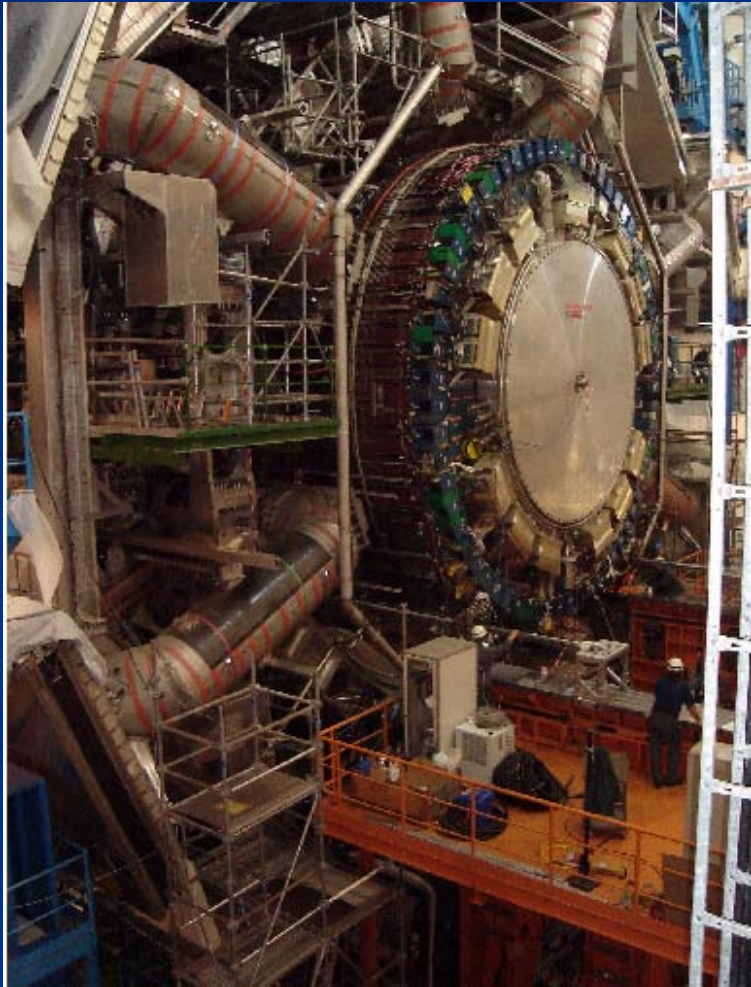
Condensed matter physics : laws depend on state of the system

- Ground state , thermal equilibrium state ...
- Example : Laws of electromagnetism in superconductor are different from Maxwells' laws

Standard model of particle physics :

Electroweak gauge symmetry is spontaneously broken by expectation value of Higgs scalar

Spontaneous symmetry breaking to be confirmed at the LHC



Cosmology :

- Universe is not in one fixed state
- Dynamical evolution
- Laws are expected to depend on time

Restoration of symmetry at high temperature in the early Universe

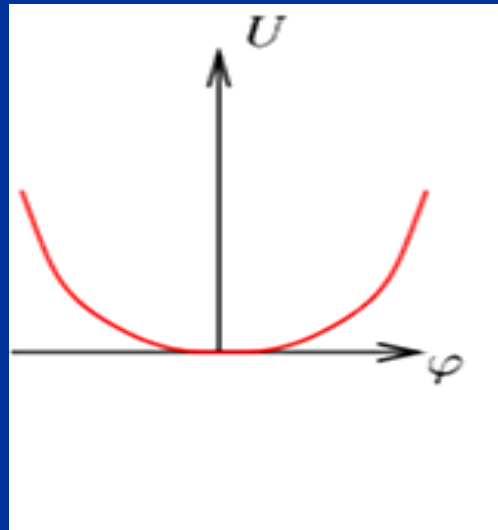
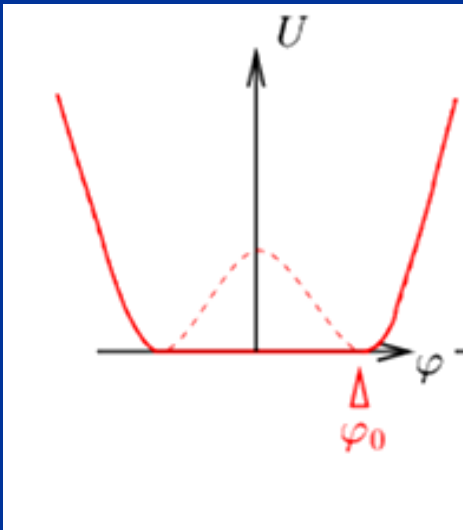
Low T
SSB

$$\langle \varphi \rangle = \varphi_0 \neq 0$$

High T
SYM

$$\langle \varphi \rangle = 0$$

high T :
Less order
More symmetry



Example:
Magnets

In hot plasma
of early Universe :

masses of electron und muon
not different!

similar strength of electromagnetic
and weak interaction

Varying couplings

only question :

How strong is **present** variation of couplings ?

Can variation of fundamental “constants” be observed ?

Fine structure constant α (electric charge)

Ratio electron mass to proton mass

Ratio nucleon mass to Planck mass

Time evolution of couplings and scalar fields

- Fine structure constant depends on value of cosmon field : $\alpha(\varphi)$

in standard model: couplings depend on value of Higgs scalar field

- Time evolution of φ  Time evolution of α

Jordan,...

Static scalar fields

In Standard Model of particle physics :

- Higgs scalar has settled to its present value around 10^{-12} seconds after big bang.
- Chiral condensate of QCD has settled at present value after quark-hadron phase transition around 10^{-6} seconds after big bang .
- No scalar with mass below pion mass.
- No substantial change of couplings after QCD phase transition.
- Coupling constants are frozen.

**Observation of time- or space-
variation of couplings**



Physics beyond Standard Model

Particle masses in quintessence cosmology

can depend on value of cosmon field

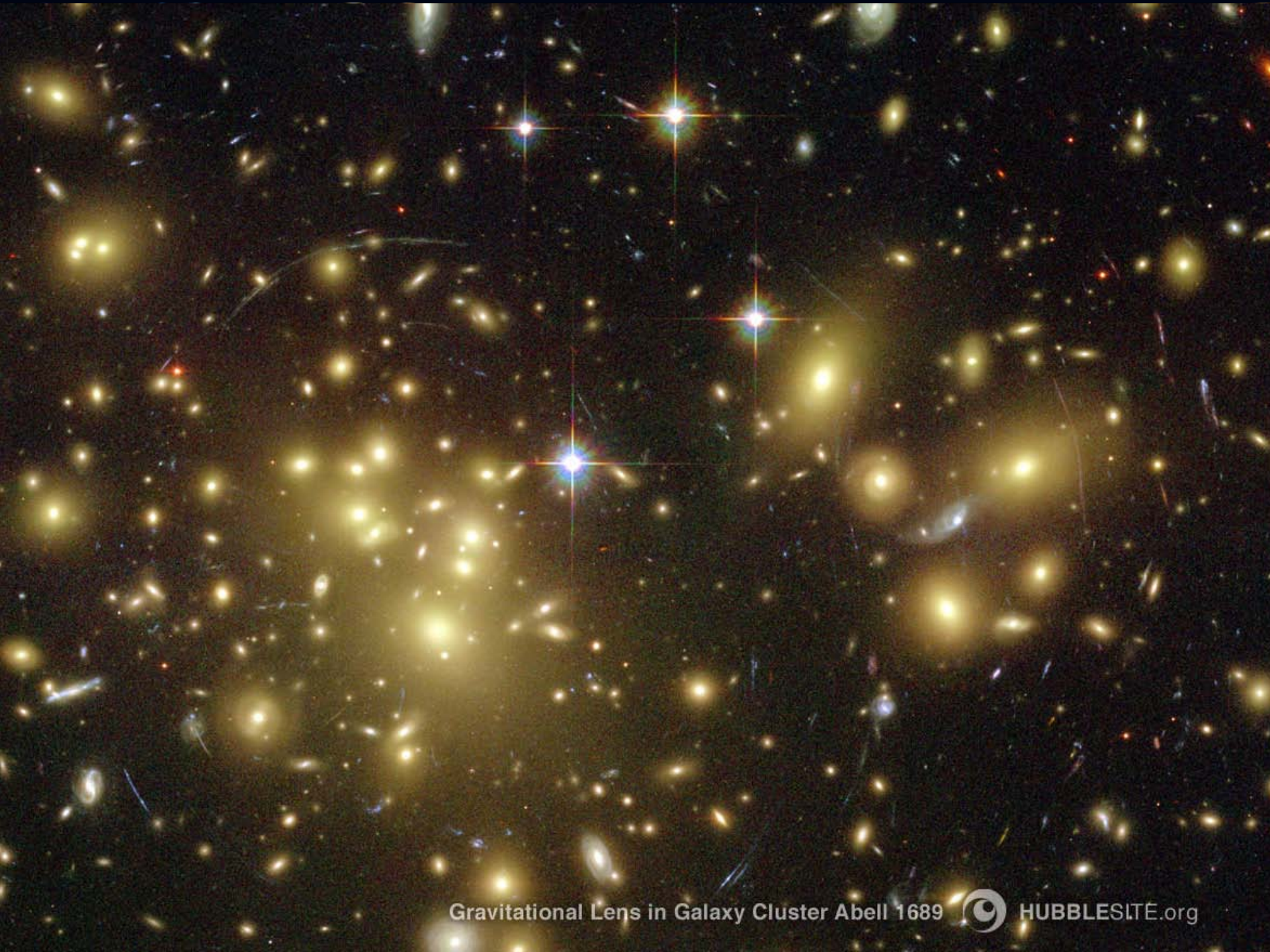
similar to dependence on value of Higgs field



Dark Energy :

Energy density that does not clump

Photons , gravitons : insignificant



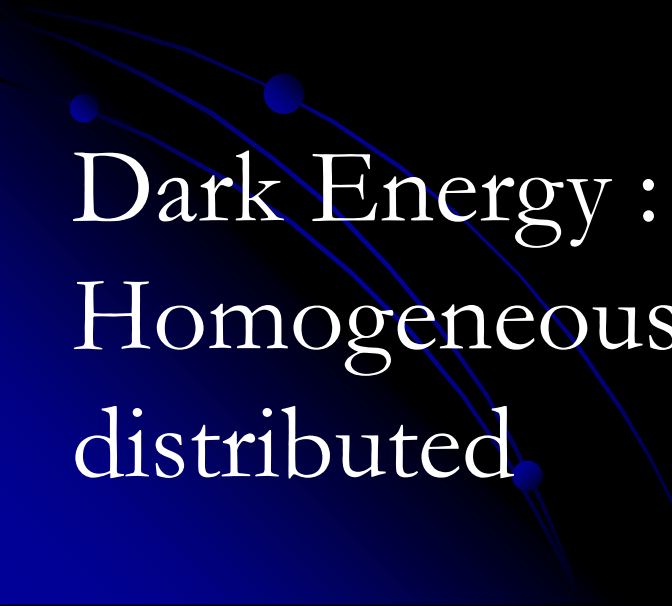
Gravitational Lens in Galaxy Cluster Abell 1689



HUBBLESITE.org

**Space between clumps
is not empty :**

Dark Energy !



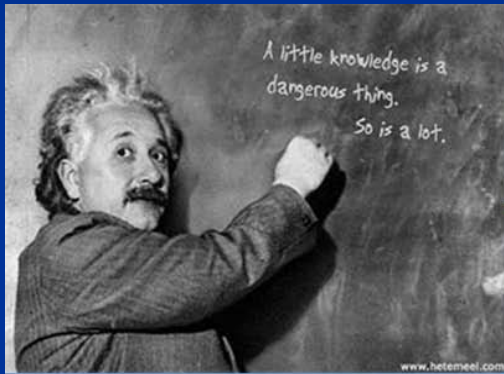
Dark Energy :
Homogeneously
distributed.

**Dark Energy density is
the same at every point of space**

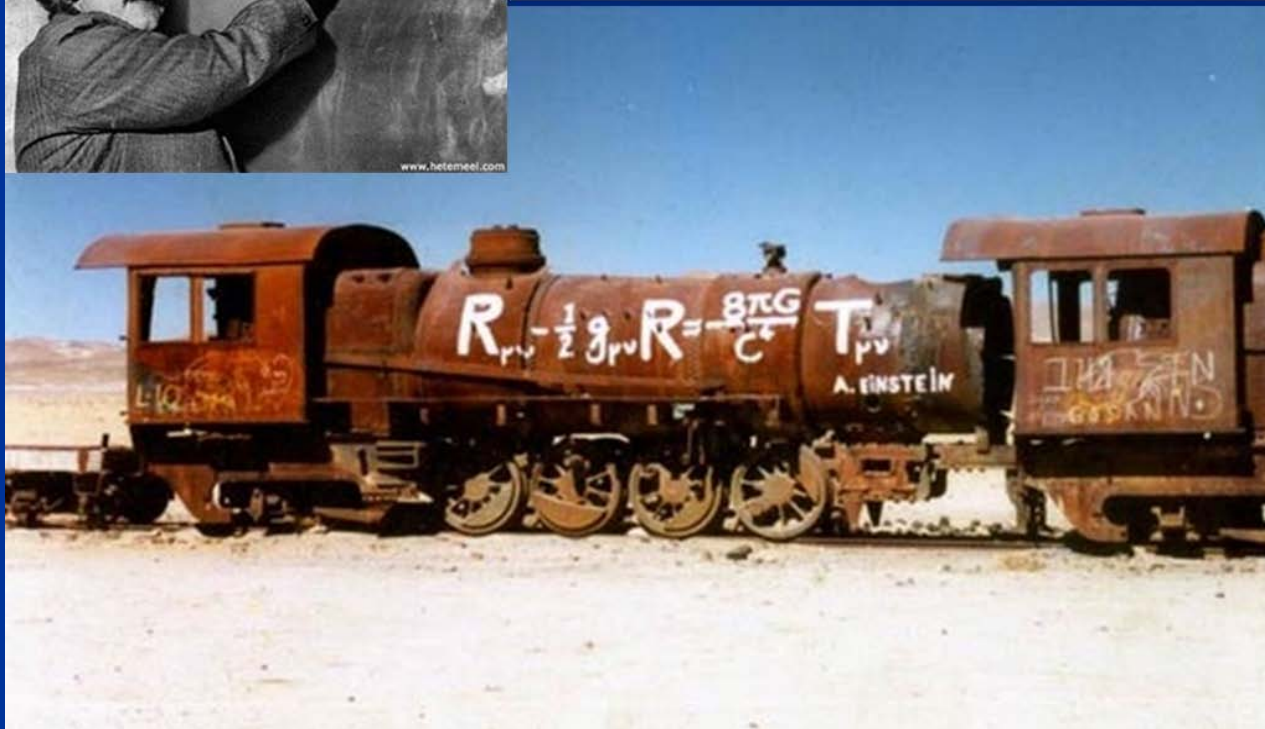
“ homogeneous “

**No force in absence of matter –
“ In what direction should it draw ? “**

Einstein's equations : almost static Dark Energy predicts accelerated expansion of Universe



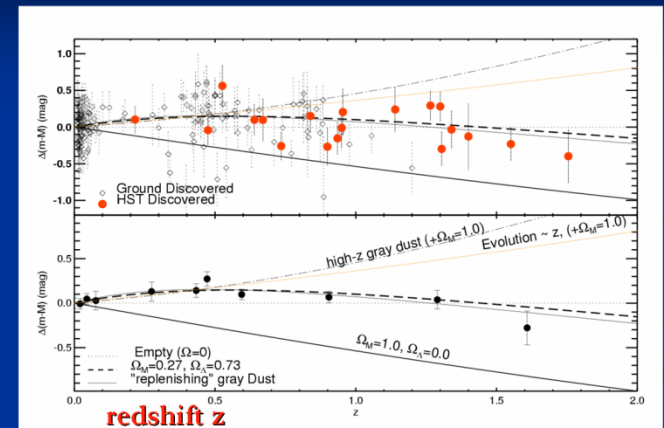
$$G_{\mu\nu} = 8\pi T_{\mu\nu}$$



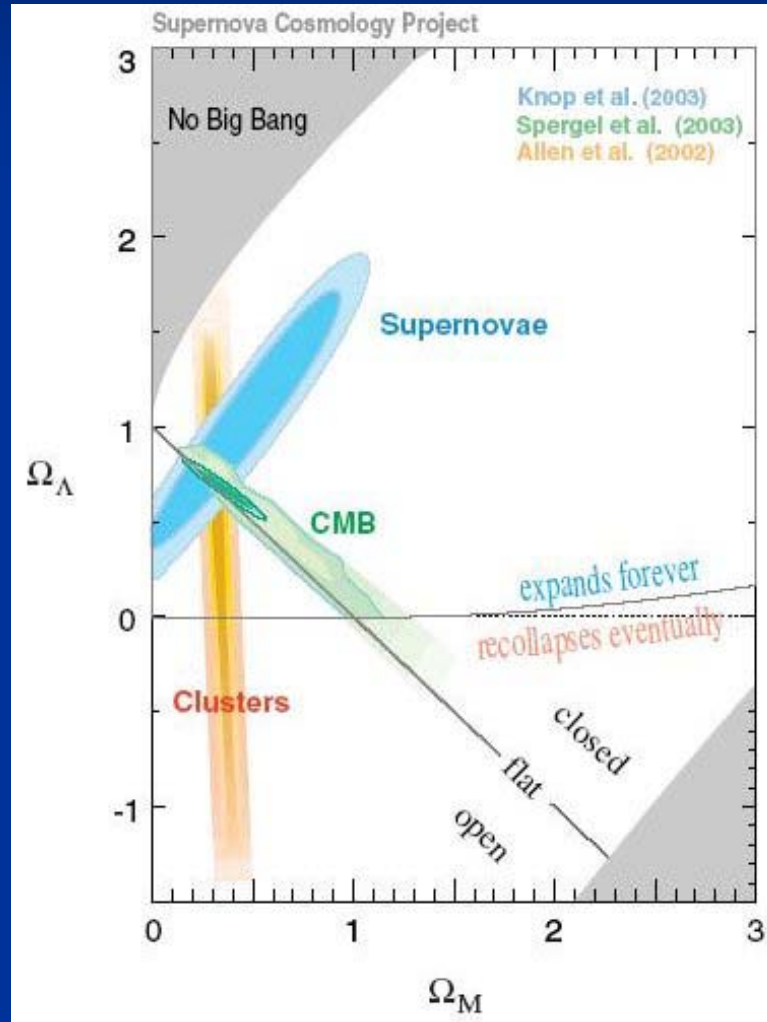
Predictions for dark energy cosmologies

*The expansion of the Universe
accelerates today !*

Supernovae 1a Hubble diagram



Dark Energy : observations fit together !



Composition of the Universe

$$\Omega_b = 0.045$$

visible

clumping

$$\Omega_{dm} = 0.2$$

invisible

clumping

$$\Omega_h = 0.75$$

invisible

homogeneous

What is Dark Energy ?

Cosmological Constant

or

Quintessence ?

Cosmological Constant

- Einstein -

- Constant λ compatible with all symmetries
- No time variation in contribution to energy density
- Why so small ? $\lambda/M^4 = 10^{-120}$
- Why important just today ?

Cosmological mass scales

- Energy density

$$\rho \sim (2.4 \times 10^{-3} \text{ eV})^{-4}$$

- Reduced Planck mass

$$M = 2.44 \times 10^{27} \text{ eV}$$

- Newton's constant

$$G_N = (8\pi M^2)$$

Only ratios of mass scales are observable !

homogeneous dark energy: $\rho_h/M^4 = 6.5 \cdot 10^{-121}$

matter: $\rho_m/M^4 = 3.5 \cdot 10^{-121}$

Time evolution

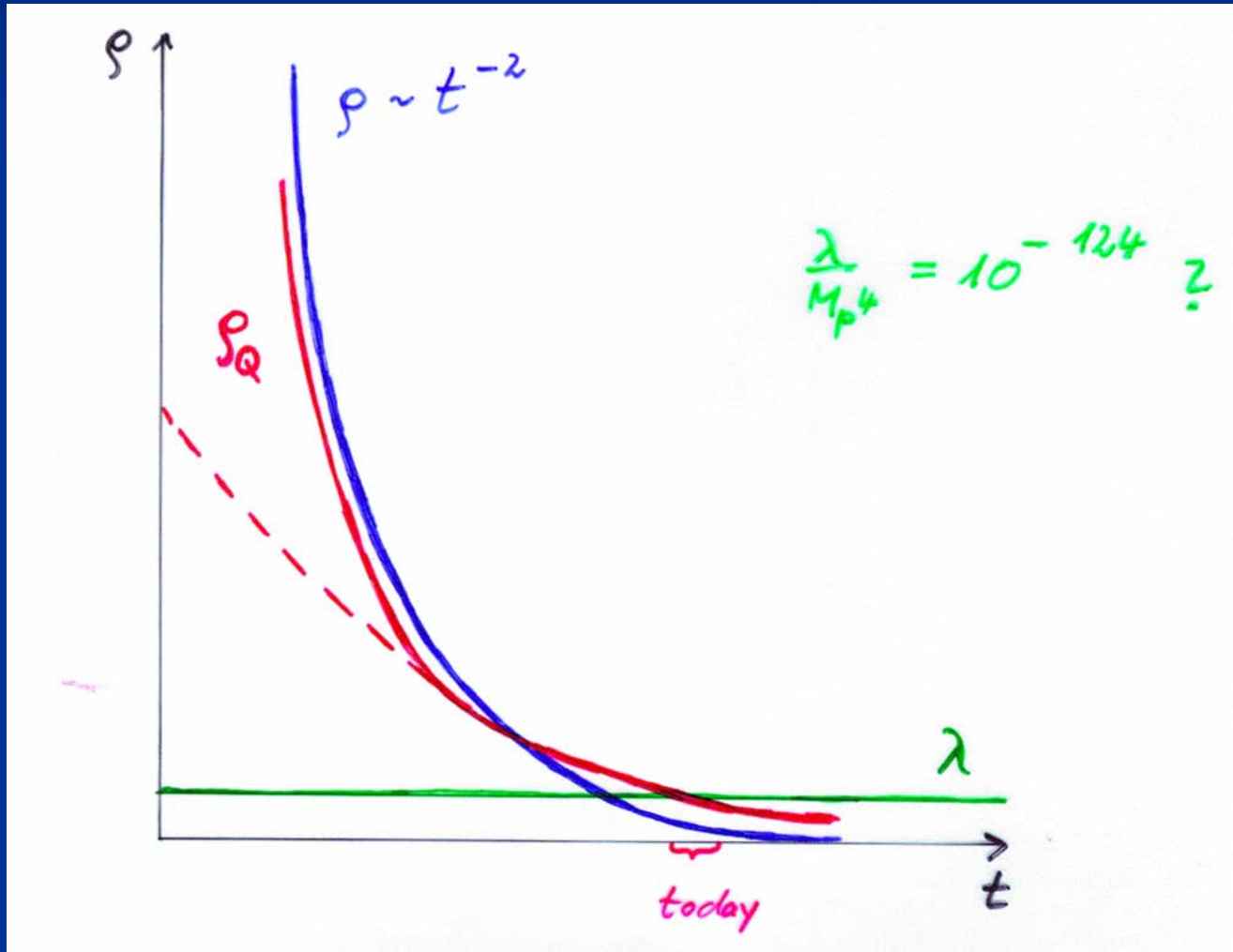
- $\rho_m/M^4 \sim a^{-3} \sim t^{-2}$ matter dominated universe
- $\rho_r/M^4 \sim a^{-4} \sim t^{-3/2}$ radiation dominated universe
- $\rho_r/M^4 \sim a^{-4} \sim t^{-2}$ radiation dominated universe

Huge age \Rightarrow small ratio

Same explanation for small dark energy?

Cosm. Const.
static

Quintessence
dynamical



Quintessence

Dynamical dark energy ,
generated by scalar field
(cosmon)

C.Wetterich,Nucl.Phys.B302(1988)668, 24.9.87
P.J.E.Peebles,B.Ratra,ApJ.Lett.325(1988)L17, 20.10.87

Prediction :

**homogeneous dark energy
influences recent cosmology**

- of same order as dark matter -

Original models do not fit the present observations
.... modifications

Quintessence

Cosmon – Field $\varphi(x,y,z,t)$

similar to electric field , but no direction (scalar field)

Homogeneous und isotropic Universe : $\varphi(x,y,z,t) = \varphi(t)$

Potential und kinetic energy of the cosmon -field
contribute to a dynamical energy density of the Universe !

Cosmon

- *Scalar field changes its value even in the **present** cosmological epoch*
- *Potential und kinetic energy of cosmon contribute to the energy density of the Universe*
- *Time - variable dark energy :
 $\rho_b(t)$ decreases with time !*

Evolution of cosmon field

Field equations

$$\ddot{\phi} + 3H\dot{\phi} = -dV/d\phi$$

$$3M^2H^2 = V + \frac{1}{2}\dot{\phi}^2 + \rho$$

Potential $V(\varphi)$ determines details of the model

$$V(\varphi) = M^4 \exp(-\alpha\varphi/M)$$

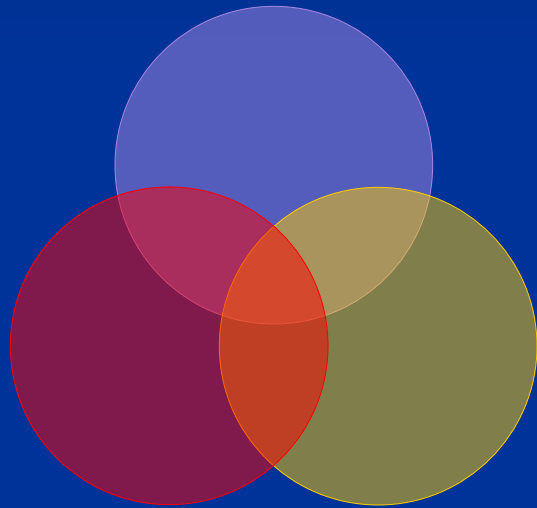
for increasing φ the potential decreases
towards zero !

Cosmon

- *Tiny mass*
- $m_c \sim H$ (depends on time !)
- *New long - range interaction*

“Fundamental” Interactions

Strong, electromagnetic, weak interactions



gravitation

cosmodynamics

On astronomical length scales:

graviton

+

cosmon

Time varying constants

- It is not difficult to obtain quintessence potentials from higher dimensional or string theories
- Exponential form rather generic
(after Weyl scaling)
- But most models show too strong time dependence of constants !

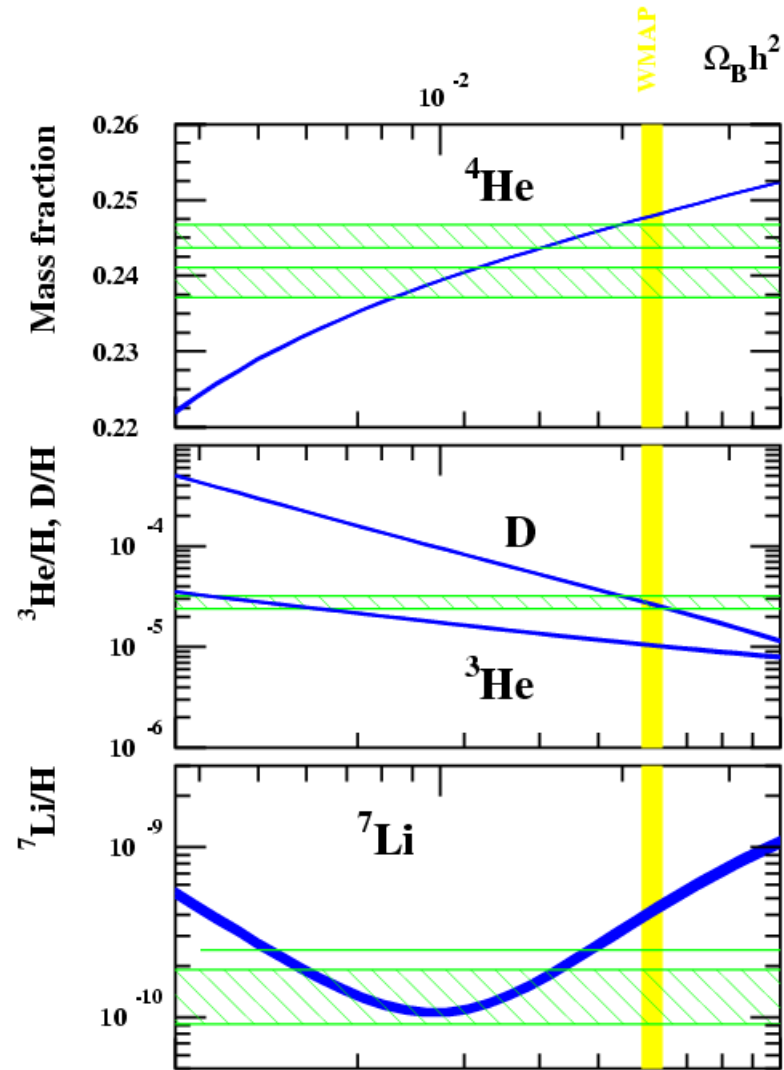
Bounds on time varying couplings from nucleosynthesis

baryons :

the matter of stars and humans

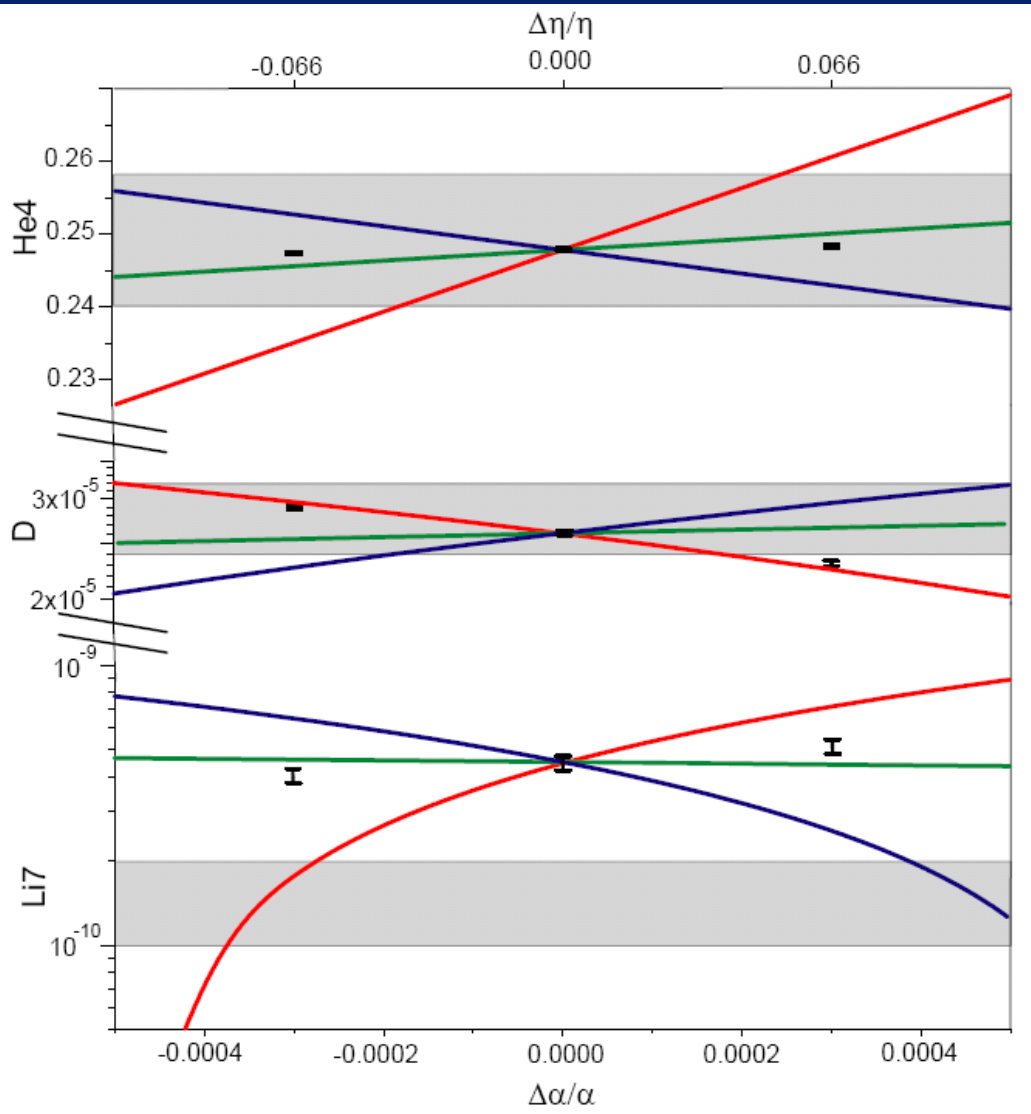
$$\Omega_b = 0.045$$

Abundancies of
primordial
light elements
from
nucleosynthesis



primordial abundances for three GUT models

He



D

Li

present
observations :
 1σ

T.Dent,
S.Stern,...

three GUT models

- unification scale \sim Planck scale
- 1) All particle physics scales $\sim \Lambda_{\text{QCD}}$
- 2) Fermi scale and fermion masses \sim unification scale
- 3) Fermi scale varies more rapidly than Λ_{QCD}

$\Delta\alpha/\alpha \approx 4 \cdot 10^{-4}$ allowed for GUT 1 and 3, larger for GUT 2

$\Delta\ln(M_n/M_p) \approx 40 \Delta\alpha/\alpha \approx 0.015$ allowed

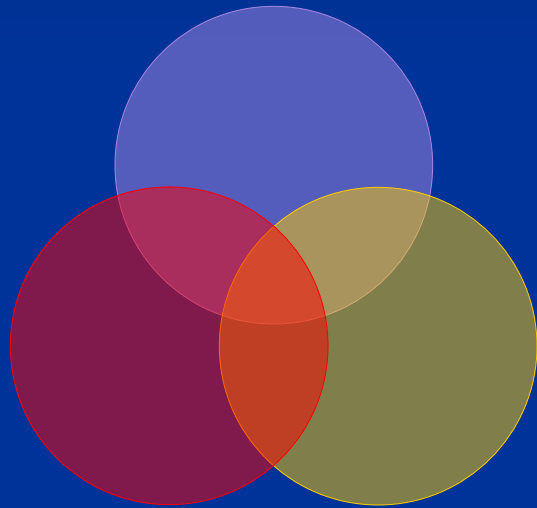
Time variation of coupling constants
must be tiny –

would be of very high significance !

Possible signal for Quintessence

“Fundamental” Interactions

Strong, electromagnetic, weak interactions



gravitation

cosmodynamics

On astronomical length scales:

graviton

+

cosmon

“Fifth Force”

- Mediated by scalar field

R.Peccei,J.Sola,C.Wetterich,Phys.Lett.B195,183(1987)

- Coupling strength: weaker than gravity
(nonrenormalizable interactions $\sim M^{-2}$)
- Composition dependence
 → violation of equivalence principle
- Quintessence: connected to time variation of fundamental couplings

C.Wetterich , Nucl.Phys.B302,645(1988)

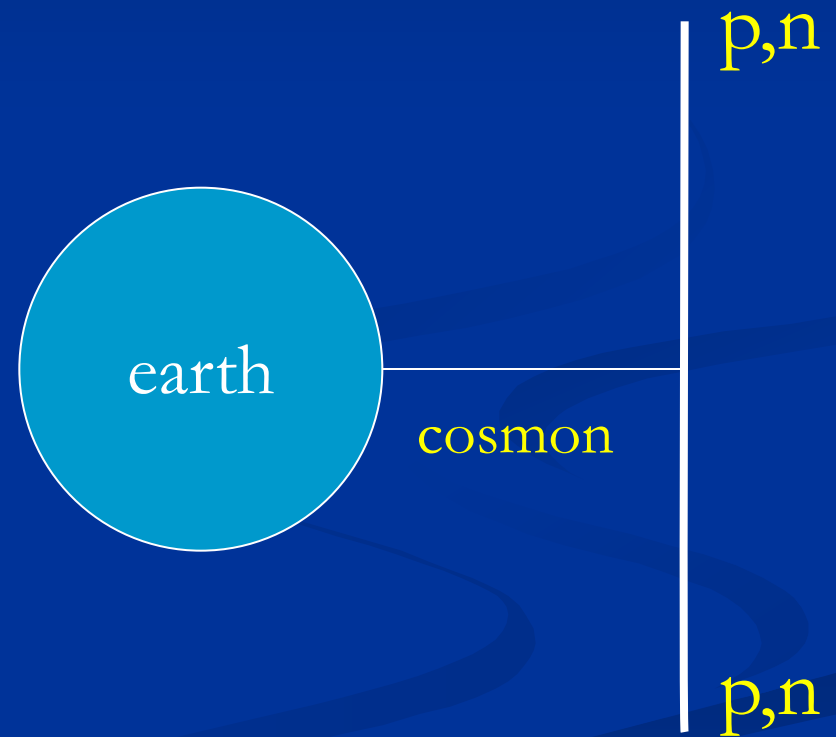
Violation of equivalence principle

Different couplings of
cosmon to proton and
neutron

Differential acceleration

“Violation of
equivalence principle”

only apparent : new “fifth force” !



Differential acceleration

Two bodies with equal mass experience
a different acceleration !

$$\eta = (a_1 - a_2) / (a_1 + a_2)$$

bound : $\eta < 3 \cdot 10^{-14}$

Cosmon coupling to atoms

- Tiny !!!
- Substantially weaker than gravity.
- Non-universal couplings bounded by tests of equivalence principle.
- Universal coupling bounded by tests of Brans-Dicke parameter ω in solar system.
- Only very small influence on cosmology.

(All this assumes validity of linear approximation)

Apparent violation of equivalence principle

and

time variation of fundamental couplings

measure both the

cosmon – coupling to ordinary matter

Differential acceleration η

For unified theories (GUT) :

$$\eta = -1.75 \cdot 10^{-2} \Delta R_z \left(\frac{\partial \ln \alpha}{\partial z} \right)^2 \frac{1 + \tilde{Q}}{\Omega_h (1 + w_h)}$$

$$\Delta R_z = \frac{\Delta Z}{Z + N} \approx 0.1$$

$$\eta = \Delta a / 2a$$

Q : time dependence of other parameters

Link between time variation of α

and violation of equivalence principle

typically : $\eta = 10^{-14}$

if time variation of α

near Oklo upper bound

to be tested (MICROSCOPE , ...)



2

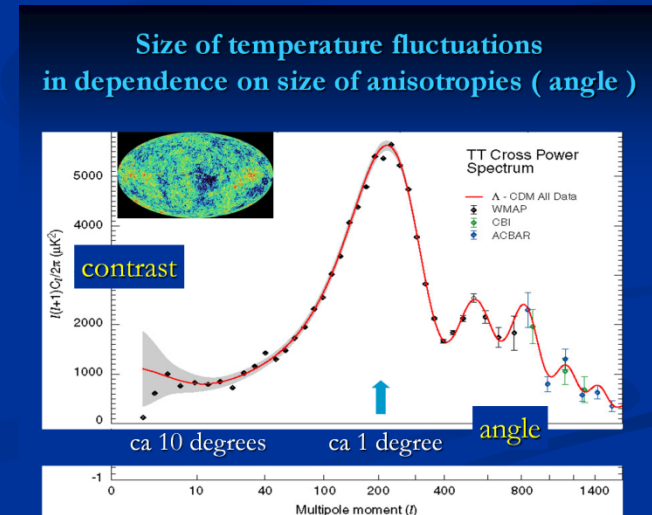
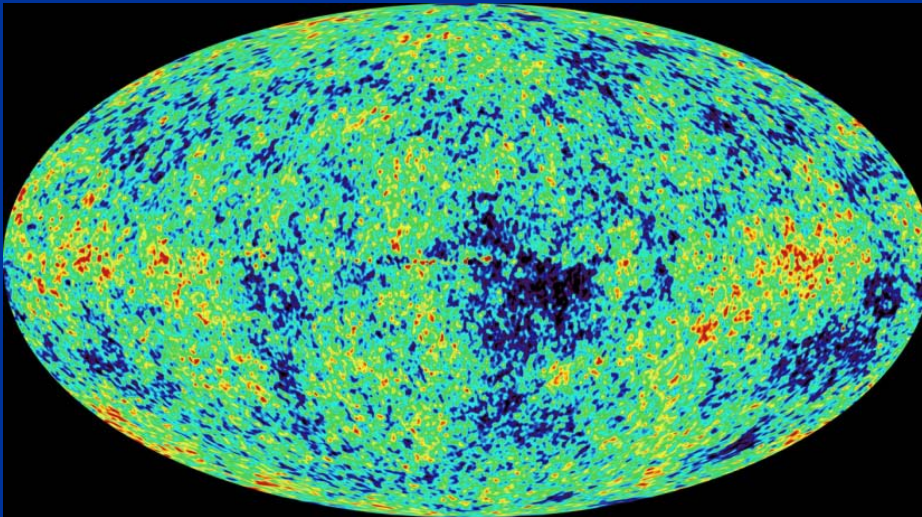
**No conflict between
quantum physics and
general relativity**

Cosmology and quantum physics

both :

Probabilistic theories

Correlations are crucial



laws are based on probabilities

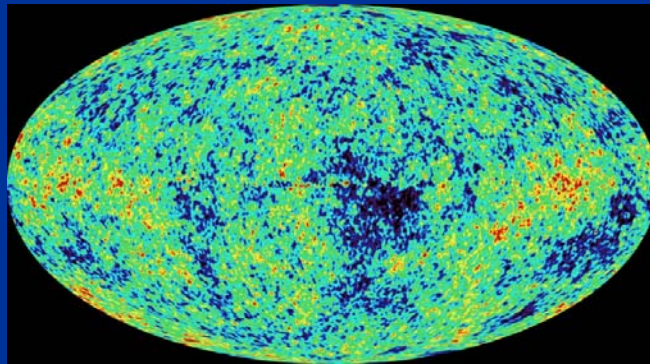
determinism as special case :
probability for event = 1 or 0

- law of big numbers
- unique ground state ...



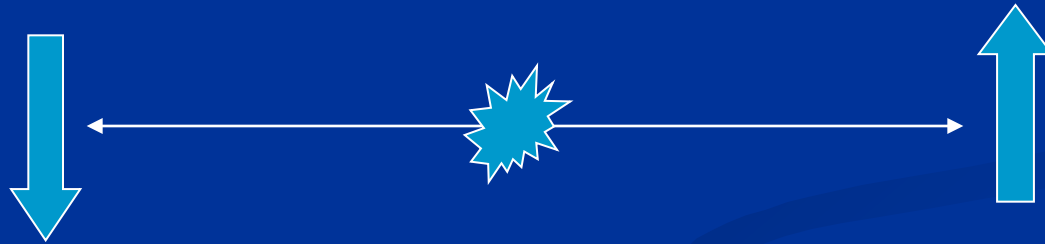
Correlations and reality

- **Correlations** are physical reality , not only the expectation values of certain observables
- Correlations can be **non-local** (also in classical statistics) ; causal processes needed only for preparation of non-local correlations at some time in the past
- Correlated subsystems cannot be separated – **the whole is more than the sum of parts**

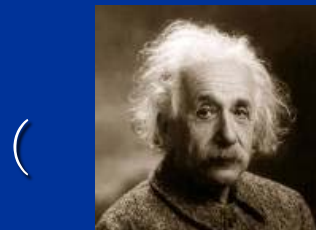


EPR - Paradoxon

Correlation between two spins is established at time of particle decay



No contradiction to causality or realism if correlations are considered as genuine part of reality



for once : not right)

conditional probability

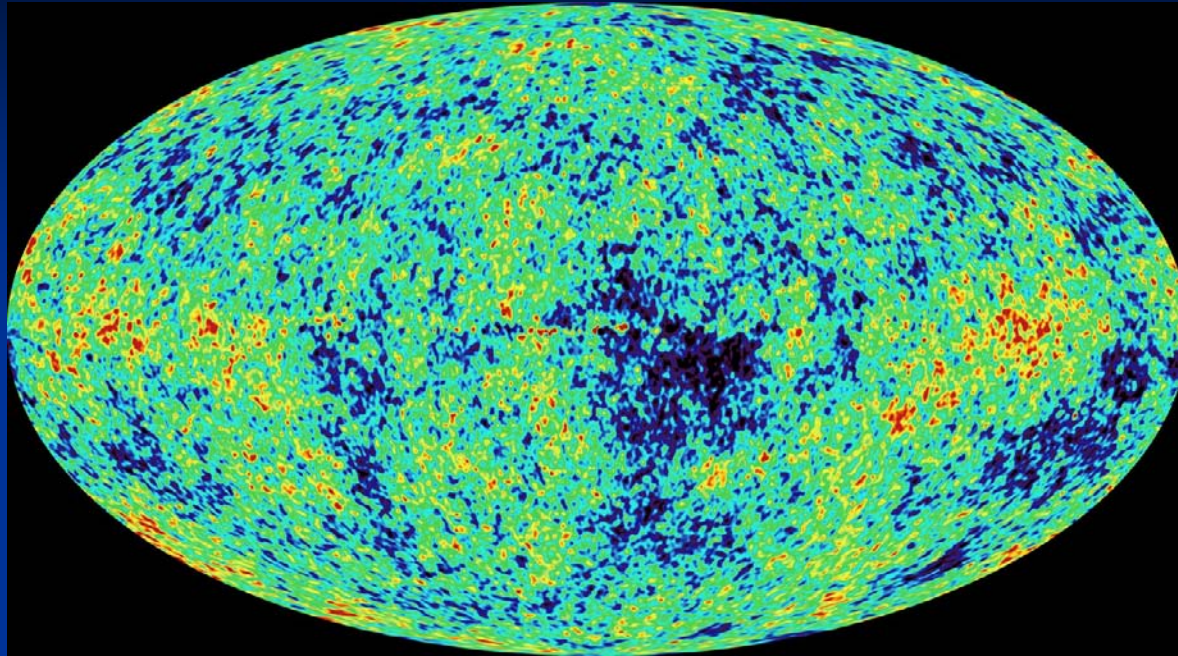
sequences of events(measurements)

are described by

conditional probabilities

*both in classical statistics
and in quantum statistics*

$w(t_1)$



not very suitable
for statement , if here and now
a pointer falls down

Schrödinger's cat

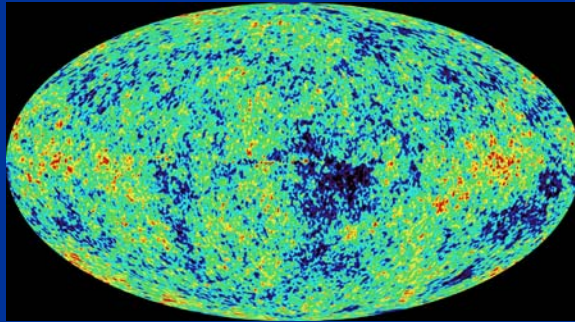


conditional probability :
if nucleus decays
then cat dead with $w_c = 1$
(reduction of wave function)

Reduction of the wave function

- Convenient way to describe conditional probabilities
- Same in classical statistics for the Universe

■ Once



realized (measured) :

new state for computation of probabilities

Metric and geometry

Quantum field theory :

- Metric is expectation value of fluctuating field
- It transforms as second rank symmetric tensor under diffeomorphisms (general coordinate transformations)
- General relativity : theory with particular symmetry – diffeomorphism symmetry
- That's all

General relativity and quantum theory

- No basic contradiction
- Metric is quantum field with same status as other fields (e.g. photon or Higgs scalar)
- Ultraviolet regularization : difficulty to implement diffeomorphism symmetry – several concepts : string theory, fixed point, lattice regularizations , lattice spinor gravity
- Metric needs not to be fundamental

Metric as collective (composite) field

Metric field could be composed from fermions

$$\tilde{g}_{\mu\nu} \sim \partial_\mu(\bar{\psi}\psi)\partial_\nu(\bar{\psi}\psi)$$

$$\bar{g}_{\mu\nu} = \langle \tilde{g}_{\mu\nu} \rangle$$

Metric as collective (composite) field

or from scalars

Metric ambiguity

given metric

$$g_{\mu\nu}$$

collective field with expectation value

$$\tilde{g}_{\mu\nu} \sim \partial_\mu(\bar{\psi}\psi)\partial_\nu(\bar{\psi}\psi) \quad \bar{g}_{\mu\nu} = \langle \tilde{g}_{\mu\nu} \rangle$$

new metric candidate

$$g_{2,\mu\nu} = \alpha g_{\mu\nu} + \beta \bar{g}_{\mu\nu}$$

Which one is physical metric ?

3

Time and space are
born at the big bang

Which metric to choose ?

Many candidates in general relativity

$$g_{2,\mu\nu} = \alpha g_{\mu\nu} + \beta R_{\mu\nu} + \gamma R g_{\mu\nu}$$

and quantum field theory

$$\tilde{g}_{\mu\nu} \sim \partial_\mu(\bar{\psi}\psi)\partial_\nu(\bar{\psi}\psi) \quad \bar{g}_{\mu\nu} = \langle \tilde{g}_{\mu\nu} \rangle$$

$$g_{2,\mu\nu} = \alpha g_{\mu\nu} + \beta \bar{g}_{\mu\nu}$$

It does not matter !

for distances large compared to Planck length

Ambiguity close to Planck scale !

no unique metric anymore close to Planck scale

different metrics – different definitions of distances

no unique geometry

Space and time in the usual sense , with definite geometry ,
emerge only after the big bang

Theories with two or several metrics

Metric potential for one metric

$$V = \lambda\sqrt{g} , \quad g = |\det g_{\mu\nu}|$$

Space-time independent metric
has to obey extremum condition

$$\frac{\partial V}{\partial g_{\mu\nu}} = \frac{1}{2}\lambda\sqrt{g}g^{\mu\nu} = 0 , \quad g^{\mu\nu}g_{\nu\rho} = \delta^{\mu}_{\rho}$$

Metric potential for two metrics

New invariants are possible

$$g_{1,\mu\nu} g_2^{\mu\nu}$$

Simple metric potential

$$\begin{aligned} V = & \alpha_1 \sqrt{g_1} + \alpha_2 \sqrt{g_2} \\ & + (\beta_1 \sqrt{g_1} + \beta_2 \sqrt{g_2}) g_{1,\mu\nu} g_2^{\mu\nu} \\ & + (\delta_1 \sqrt{g_1} + \delta_2 \sqrt{g_2}) g_{2,\mu\nu} g_1^{\mu\nu} . \end{aligned}$$

Proportionality of metrics

ansatz:

$$g_{1,\mu\nu} = \sigma_1 \eta_{\mu\nu}, \quad g_{2,\mu\nu} = \sigma_2 \eta_{\mu\nu}$$

$$\gamma = \sigma_2 / \sigma_1$$

$$V = \sigma_1^2 \left(\sum_n s_n \gamma^n \right) = \sigma_1^2 W(\gamma).$$

Extremum of $W(\gamma)$ at γ_0 with $W(\gamma_0) = W_0$

$$W_0 = 0$$



solution with

$$g_{1,\mu\nu} = \eta_{\mu\nu}, \quad g_{2,\mu\nu} = \gamma_0 \eta_{\mu\nu}$$

Two metrics are equivalent if

$$g_{2,\mu\nu} = \gamma g_{1,\mu\nu}$$

only units for coordinates differ

Massive tensors

expansion

$$g_{1,\mu\nu} = g_{\mu\nu} , \quad g_{2,\mu\nu} = \gamma(\delta_{\mu}^{\rho} + f_{\mu}^{\rho})g_{\rho\nu}$$

tensor fields

$$f = f_{\mu}^{\nu} \delta_{\nu}^{\mu}$$

$$\tilde{f}_{\mu}^{\nu} = f_{\mu}^{\nu} - \frac{1}{4} f \delta_{\mu}^{\nu} , \quad \tilde{f}_{\mu}^{\nu} \delta_{\nu}^{\mu} = 0,$$

$$V = \sqrt{g} \left\{ \lambda + \frac{1}{2} \mu f^2 + \frac{1}{4} \nu \tilde{f}_{\mu}^{\nu} \tilde{f}_{\nu}^{\mu} \right\}$$

mass terms

$$\lambda = \alpha_1 + \alpha_2 \gamma^2 + \frac{4}{\gamma} (\beta_1 + \beta_2 \gamma^2) + 4\gamma (\delta_1 + \delta_2 \gamma^2)$$

$$\mu = \frac{\beta_1}{2\gamma} + \frac{\alpha_2}{8} \gamma^2 + \frac{3\delta_2}{2} \gamma^3,$$

$$\nu = 4 \frac{\beta_1}{\gamma} - \alpha_2 \gamma^2 - 4\delta_2 \gamma^3.$$

Kinetic term

$$\begin{aligned} \mathcal{L}_{\text{kin}} = \sqrt{g} \left\{ & -\frac{M^2}{2} R + \frac{1}{2} Z \partial^\mu f \partial_\mu f \right. \\ & + \frac{1}{4} W_1 D^\rho \tilde{f}_\mu{}^\nu D_\rho \tilde{f}_\nu{}^\mu + \frac{1}{2} W_2 D_\rho \tilde{f}_\mu{}^\rho D_\sigma \tilde{f}_\nu{}^\sigma g^{\mu\nu} \\ & \left. + \frac{1}{2} W_3 \partial^\nu f D_\mu \tilde{f}_\nu{}^\mu \right\}. \end{aligned} \quad ($$

Field equation (for $W_2=0, W_3=0$)

$$(W_1 \partial^2 - \nu) \tilde{f}_\mu{}^\nu$$

Massive tensor field

$$(W_1 \partial^2 - \nu) \tilde{f}_\mu{}^\nu$$

Typical size of mass : Planck mass , $\nu/W_1 \sim M^2$

Yukawa type interaction , negligible for distances larger than Planck length

Relevant for distances around Planck length

No unique metric at Planck scale

$$g_{1,\mu\nu} = g_{\mu\nu} , \quad g_{2,\mu\nu} = \gamma(\delta_{\mu}^{\rho} + f_{\mu}^{\rho})g_{\rho\nu}$$

Proportionality between both metrics is no longer valid once f plays a role!

Energy density

$$\begin{aligned}
 \rho = & \frac{1}{2}\mu f^2 + \frac{1}{4}\nu \left\{ \sum_k (\tilde{f}_k^k)^2 + \left(\sum_k \tilde{f}_k^k \right)^2 \right\} \\
 & + \frac{1}{2}\nu \sum_k \left\{ \left(\tilde{f}_0^k \right)^2 + \sum_{l>k} (\tilde{f}_k^l)^2 \right\} \\
 & + \frac{1}{2}Z(\partial_0 f)^2 + \frac{1}{2}Z \sum_i (\partial_i f)^2 \\
 & + \frac{1}{4}W_1 \left\{ \sum_i \left[\sum_k (\partial_i \tilde{f}_k^k)^2 + (\partial_i \sum_k \tilde{f}_k^k)^2 \right. \right. \\
 & \left. \left. + 2 \sum_{k,l>k} (\partial_i \tilde{f}_k^l)^2 + 2 \sum_k (\partial_i \tilde{f}_0^k)^2 \right] \right. \\
 & \left. + \sum_k (\partial_0 \tilde{f}_k^k)^2 + (\partial_0 \sum_k \tilde{f}_k^k)^2 \right. \\
 & \left. + 2 \sum_{k,l>k} (\partial_0 \tilde{f}_k^l)^2 - 6 \sum_k (\partial_0 \tilde{f}_0^k)^2 \right\}.
 \end{aligned}$$

Conclusions

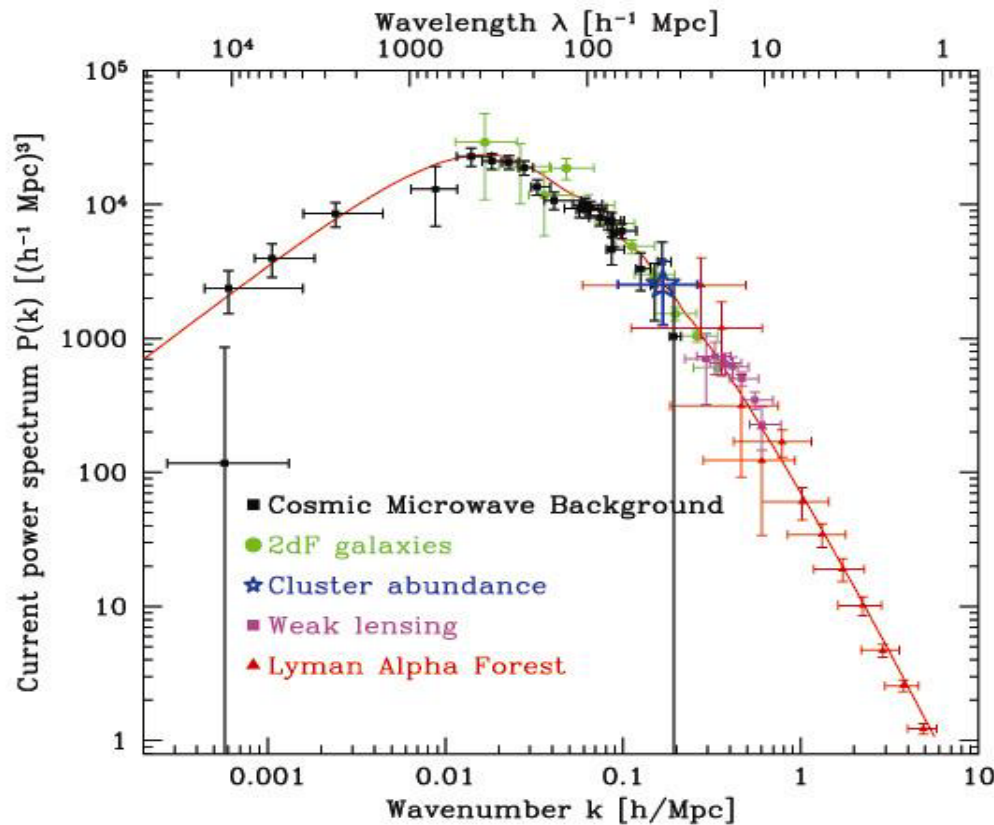
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End

Structure formation :

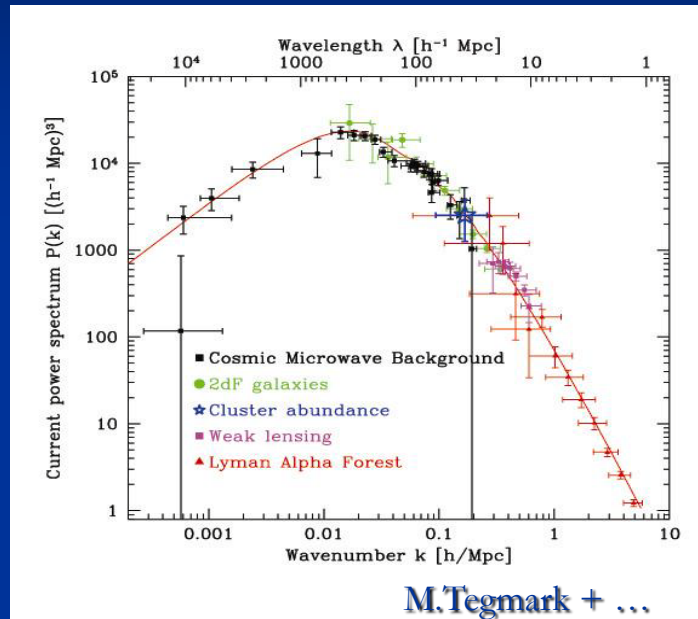
One primordial fluctuation spectrum



Waerbeke

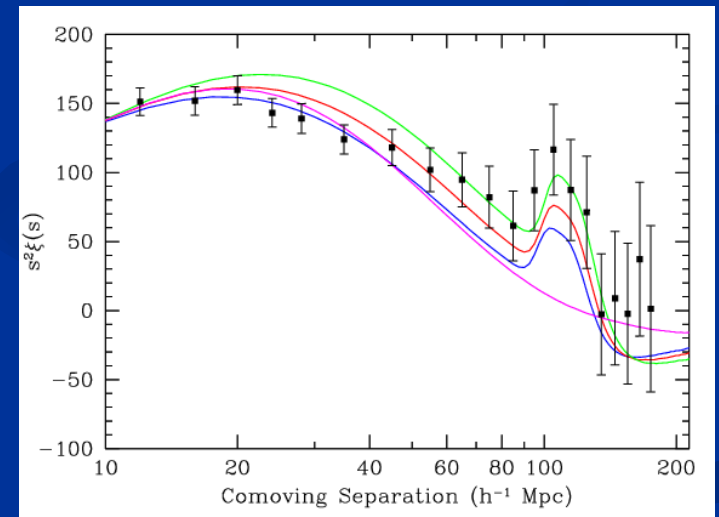
CMB agrees with
Galaxy distribution
Lyman - α
and
Gravitational
Lensing !

Power spectrum



Baryon - Peak

galaxy –
correlation –
function



Structure formation :
One primordial
fluctuation- spectrum

SDSS