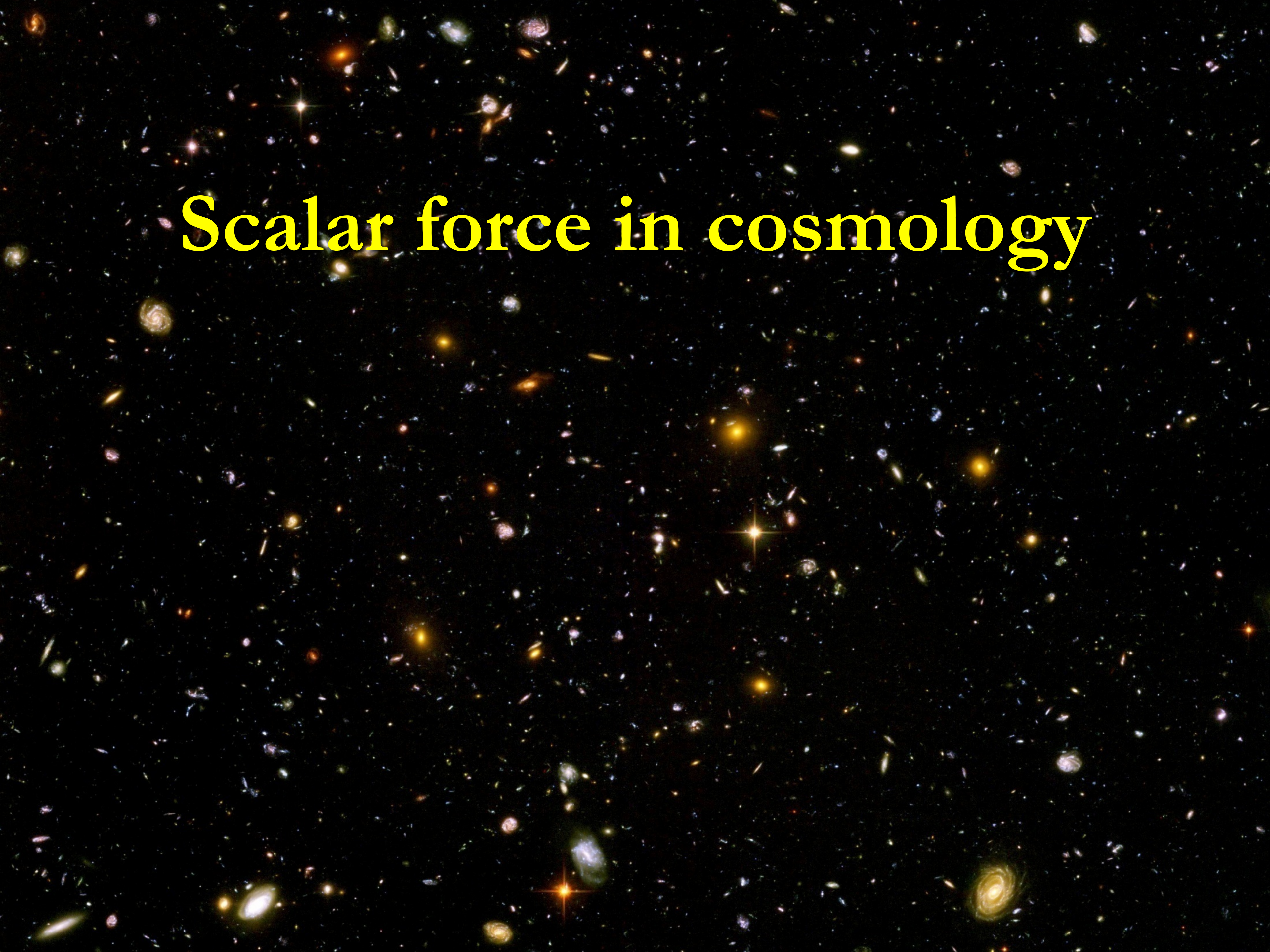
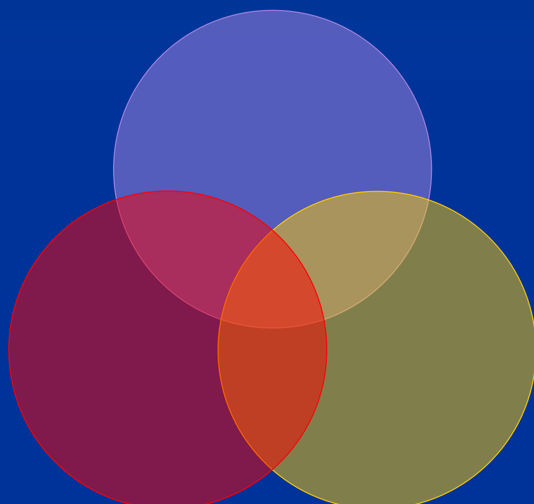
Scalar force in cosmology



"Fundamental" Interactions

Strong, electromagnetic, weak interactions



gravitation cosmodynamics

On astronomical length scales:

graviton

+

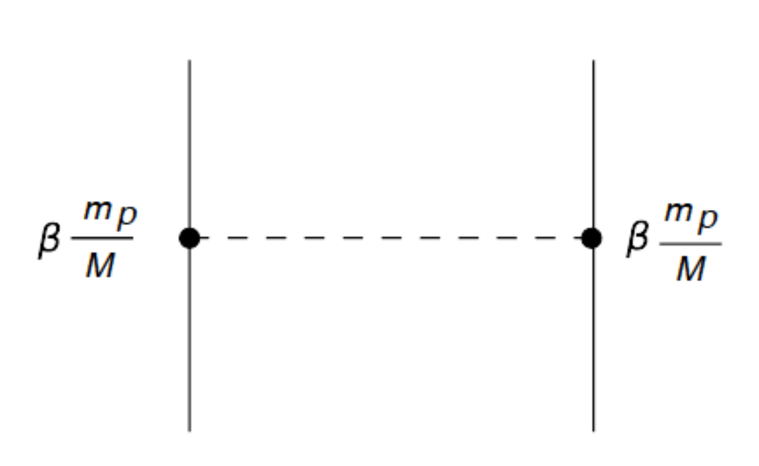
cosmon

(A) Fifth force

Scalar tensor theories

- Dirac's hypothesis
- Weyl scaling
- Brans-Dicke theory

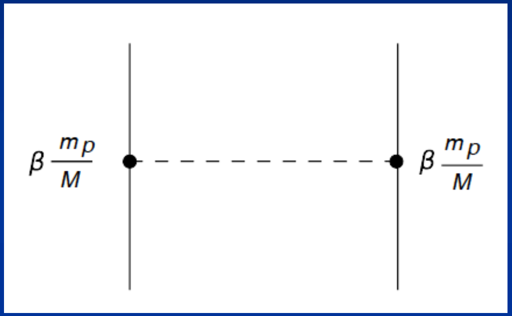
additional interaction :



Fifth force

- coupling strength $\beta \frac{m_p}{M}$

- range $\frac{1}{m}$



m ~ H : long range (dynamical dark energy)

m $\gg$ H : short range (dark matter)

concentrate on long range – effectively massless scalar for earth, solar system, galaxies, black holes, laboratory

Models with massless scalar fields

- Higher dimensions : strings, Kaluza-Klein theories (moduli)
- Modified gravity (Brans-Dicke theory, f(R)-theories, ….
- if scalar plays a role for late cosmology : “ cosmon “

Standard model with non-renormalizable operators

$$\beta_m \frac{\varphi}{M} \bar{\psi}\psi h$$

φ - dependent masses

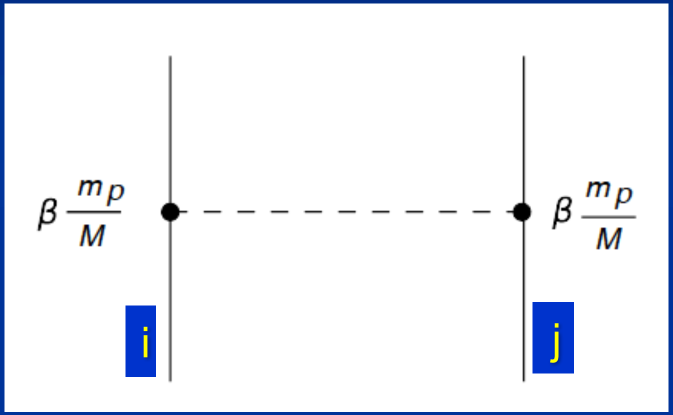
$$\beta_F \frac{\varphi}{M} F_{\mu\nu} F^{\mu\nu}$$

φ - dependent gauge couplings

Tests of GR : $|\beta_n| \ll 10^{-2}$

Tests of equivalence principle: $|\Delta\beta_i| \ll 10^{-5}$

Scalar force (fifth force)



attractive: $\beta_i \beta_j > 0$

repulsive: $\beta_i \beta_j < 0$

“Fifth Force”

- Mediated by scalar field

R.Peccei,J.Sola,C.Wetterich,Phys.Lett.B195,183(1987)

- Coupling strength: weaker than gravity

(non-renormalizable interactions ~ M^{-2})

- Composition dependence

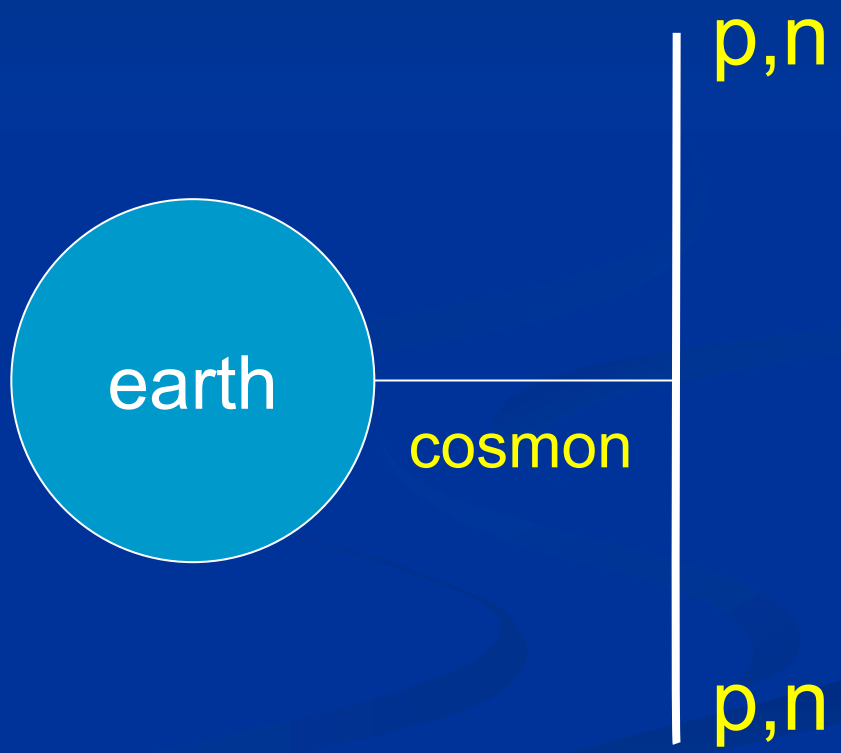
⇨ violation of equivalence principle

- Quintessence: connected to time variation of fundamental couplings

C.Wetterich , Nucl.Phys.B302,645(1988)

Violation of equivalence principle

- Different couplings of cosmon to proton and neutron
- Differential acceleration
- “Violation of equivalence principle”



only apparent : new “fifth force” !

Differential acceleration

Two bodies with equal mass experience
a different acceleration !

$$\eta = (a_1 - a_2) / (a_1 + a_2)$$

bound : $\eta < 3\ 10^{-14}$

Fundamental "constants" are not constant

Have coupling constants in the early Universe other values than today ?

Yes !

Fundamental couplings in quantum field theory

Masses and coupling constants are determined by properties of ***vacuum*** *!*

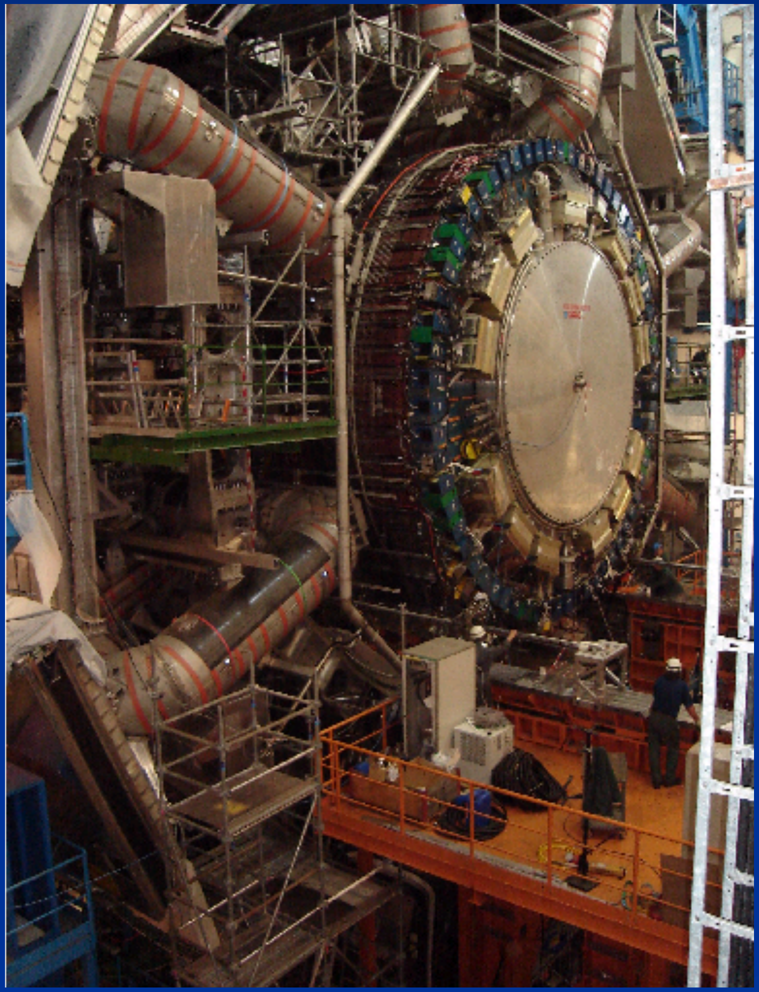
Similar to Maxwell – equations in matter

Standard model of particle physics :

Electroweak gauge symmetry is spontaneously broken by expectation value of Higgs scalar

Quark and lepton masses proportional to value of Higgs scalar

Spontaneous symmetry breaking confirmed at the LHC



Cosmology :

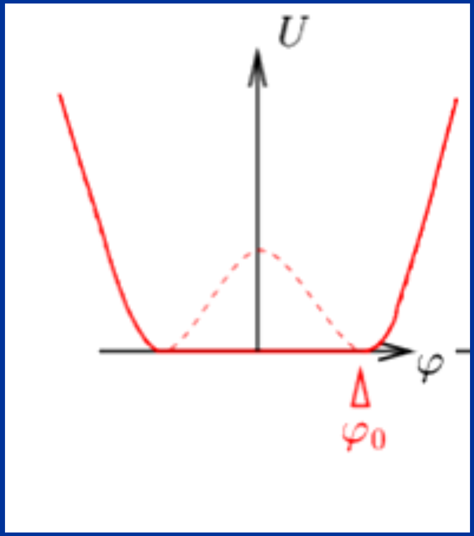
- Universe is not in one fixed state
- Dynamical evolution
- Laws are expected to depend on time

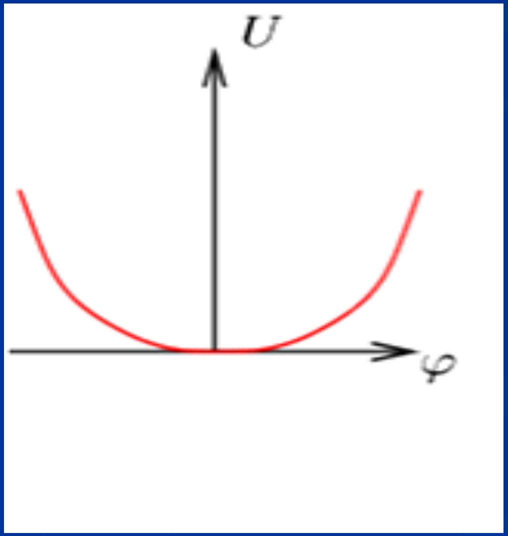
Restoration of symmetry at high temperature in the early Universe

Low T
SSB
$<\varphi>=\varphi_0 \neq 0$

High T
SYM
$<\varphi>=0$

high T :
Less order
More symmetry





Example:
Magnets

In hot plasma of early Universe :

masses of electron und muon not different!

similar strength of electromagnetic and weak interaction

Varying couplings

only question :

How strong is **present** variation of couplings ?

Can variation of fundamental "constants" be observed ?

Fine structure constant α (electric charge)

Ratio electron mass to proton mass

Ratio nucleon mass to Planck mass

Time evolution of couplings and scalar fields

- Fine structure constant depends on value of scalar field : $\alpha(\varphi)$

- Time evolution of φ ⟹

  Time evolution of α

Jordan,…

Static scalar fields

In Standard Model of particle physics :

- Higgs scalar has settled to its present value around 10^{-12} seconds after big bang.
- Chiral condensate of QCD has settled at present value after quark-hadron phase transition around 10^{-6} seconds after big bang .
- No scalar with mass below pion mass.
- No substantial change of couplings after QCD phase transition.
- Coupling constants are frozen.

Observation of time- or space- variation of couplings

→

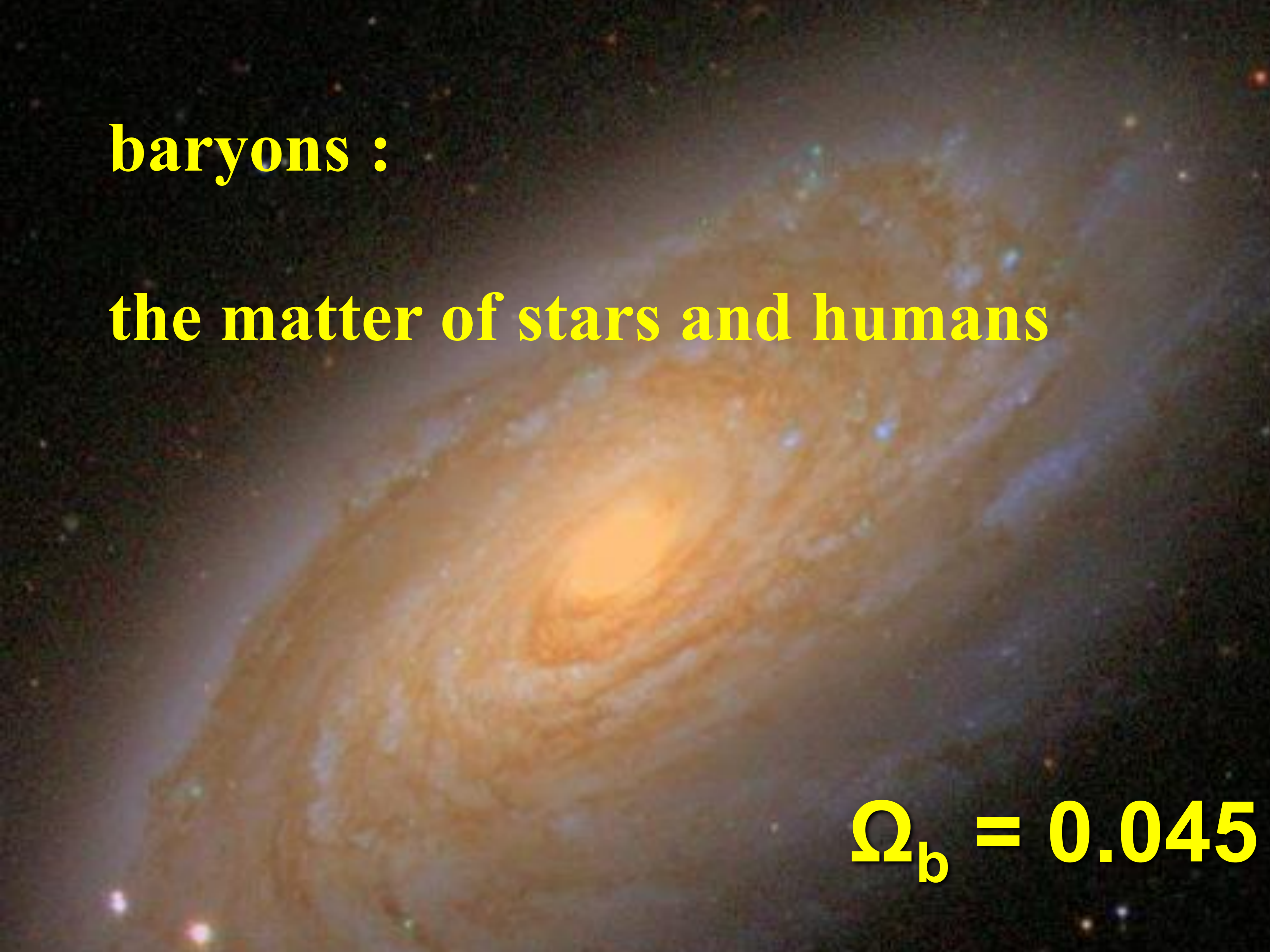
Physics beyond Standard Model

Particle mass ratios and dimensionless couplings in quintessence cosmology

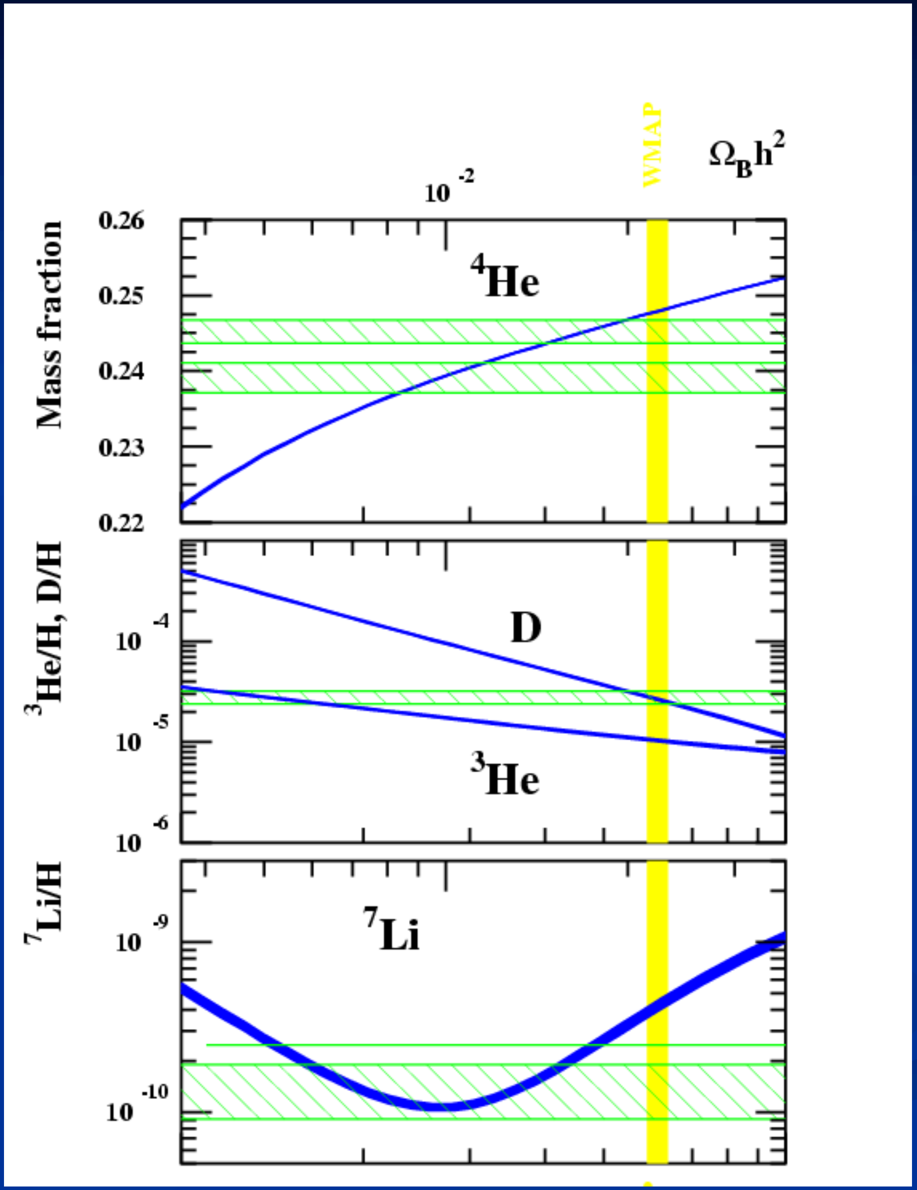
can depend on value of cosmon field

similar to dependence on value of Higgs field

Bounds on time varying couplings from nucleosynthesis


baryons :
the matter of stars and humans
$\Omega_b = 0.045$

Abundancies of primordial light elements from nucleosynthesis



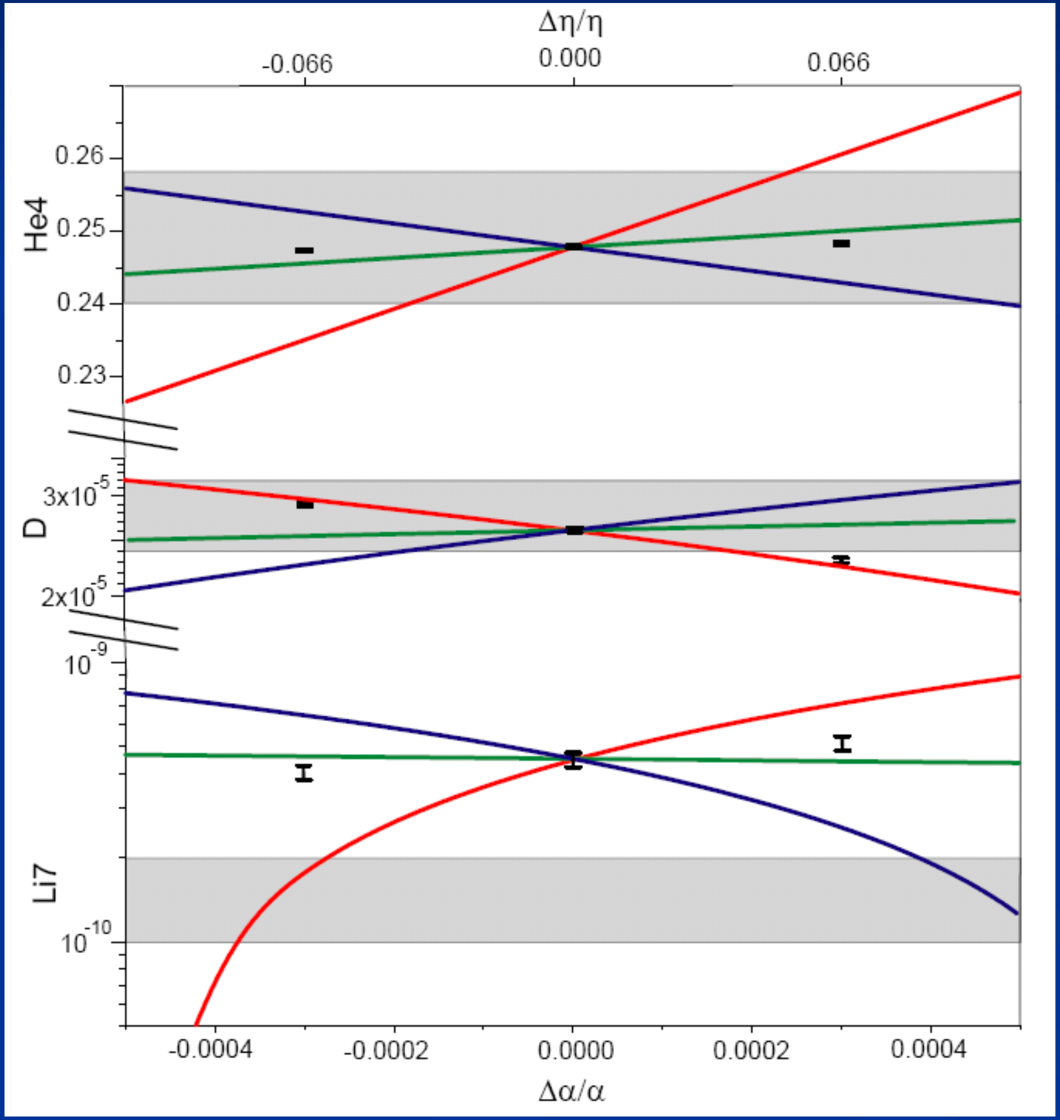
A.Coc

primordial abundances for three GUT models

He

D

Li



present observations : 1σ

T.Dent, S.Stern,…

three GUT models

- unification scale ~ Planck scale
- 1) All particle physics scales $\sim\Lambda_{QCD}$
- 2) Fermi scale and fermion masses ~ unification scale
- 3) Fermi scale varies more rapidly than Λ_{QCD}

$\Delta\alpha/\alpha \approx 4\ 10^{-4}$ allowed for GUT 2 and 3 , larger for GUT 1

$\Delta \ln(M_n/M_P) \approx 40\ \Delta\alpha/\alpha \approx 0.015$ allowed

Time variation of coupling constants must be tiny –

would be of very high significance !

Possible signal for Quintessence

Apparent violation of equivalence principle

and

time variation of fundamental couplings

measure both the

cosmon – coupling to ordinary matter

Differential acceleration η

For unified theories (GUT) :

$$\eta = -1.75\ 10^{-2} \Delta R_z \left(\frac{\partial \ln \alpha}{\partial z}\right)^2 \frac{1+\tilde{Q}}{\Omega_h (1+w_h)}$$

$$\Delta R_z = \frac{\Delta Z}{Z+N} \approx 0.1$$

$\eta = \Delta a / 2a$

Q : time dependence of other parameters

Link between time variation of α

and violation of equivalence principle

typically : $\eta = 10^{-14}$

if time variation of α

near Oklo upper bound

to be tested (MICROSCOPE , …)

Cosmon coupling to atoms

- Tiny !!!
- Substantially weaker than gravity.
- Non-universal couplings bounded by tests of equivalence principle.
- Universal coupling bounded by tests of Brans-Dicke parameter ω in solar system.
- Only very small influence on cosmology.

(All this assumes validity of linear approximation)

Key questions (1)

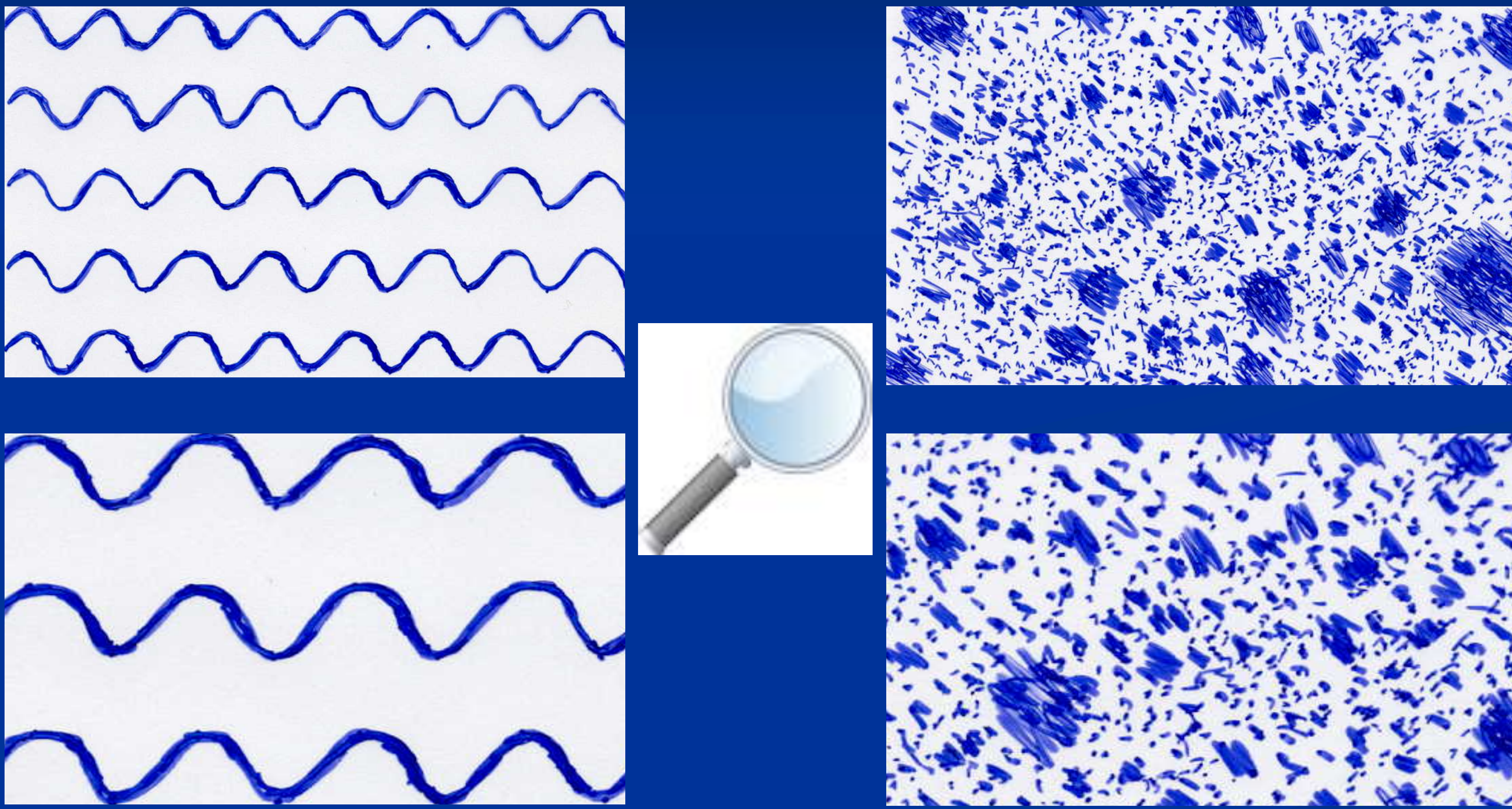
- Why is atom – scalar coupling much weaker than gravity?
- Why is scalar (almost) massless despite the presence of quantum fluctuations ?

(B) Spontaneously broken scale symmetry

scale symmetry

No intrinsic length or mass scale present

Scale symmetry



no scale symmetry

scale symmetry

Scale symmetry

All parameters with mass are proportional to scalar field χ

- Dimensionless couplings are independent of χ.
- All particle masses are proportional to χ .
- Ratios of particle masses remain constant.
- Compatibility with observational bounds on time dependence of particle mass ratios.

Scale symmetry

All parameters with mass are proportional to scalar field χ

gravity:

$$\mathcal{L} = -\frac{1}{2}\chi^2 R + \frac{B-6}{2}\partial^\mu\chi\partial_\mu\chi + \lambda\chi^4$$

Fujii, Zee

Higgs vev,

electron mass

$$\langle h \rangle \sim \chi$$

$$m_e \sim \chi$$

QCD,

proton mass

$$\Lambda_{QDC} \sim \chi$$

$$m_p \sim \chi$$

Spontaneous breaking of scale symmetry

- expectation value of scalar field breaks scale symmetry spontaneously
- massive particles are compatible with scale symmetry
- in presence of massive particles : sign of exact scale symmetry is exactly massless Goldstone boson – the dilaton

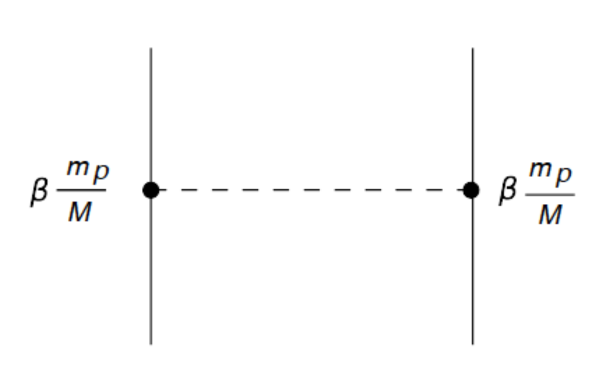
Einstein frame

Weyl scaling :

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu} \; , \; \varphi = \frac{2M}{\alpha} \ln\left(\frac{\chi}{\mu}\right)$$

scale transformation :

$$\varphi \to \varphi + \text{const.}$$



$$m = 0, \;\; \beta_i = 0$$

Key questions (1)

- Why is atom – scalar coupling much weaker than gravity?
- Why is scalar (almost) massless despite the presence of quantum fluctuations ?

answer

Scale symmetry explains both massless scalar field and vanishing couplings

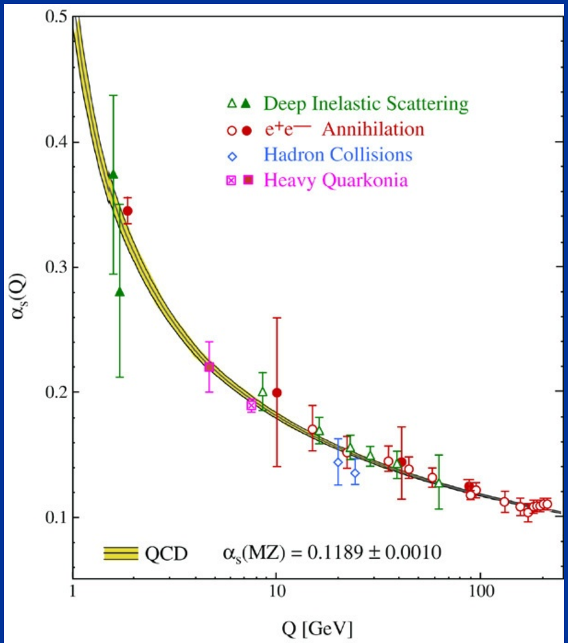
Key question (2)

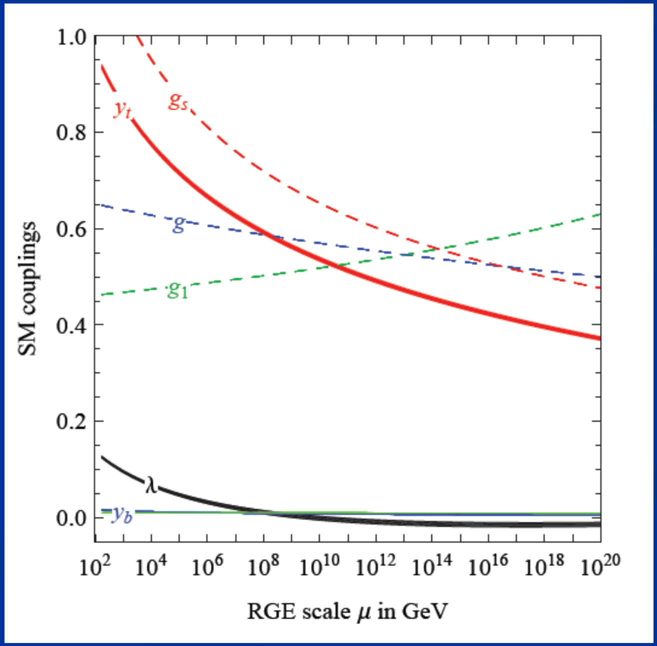
- Why does scale symmetry play a role even though it is violated by quantum fluctuations ?

quantum gravity with scalar field – the role of scale symmetry

fluctuations induce running couplings

- violation of scale symmetry
- well known in QCD or standard model

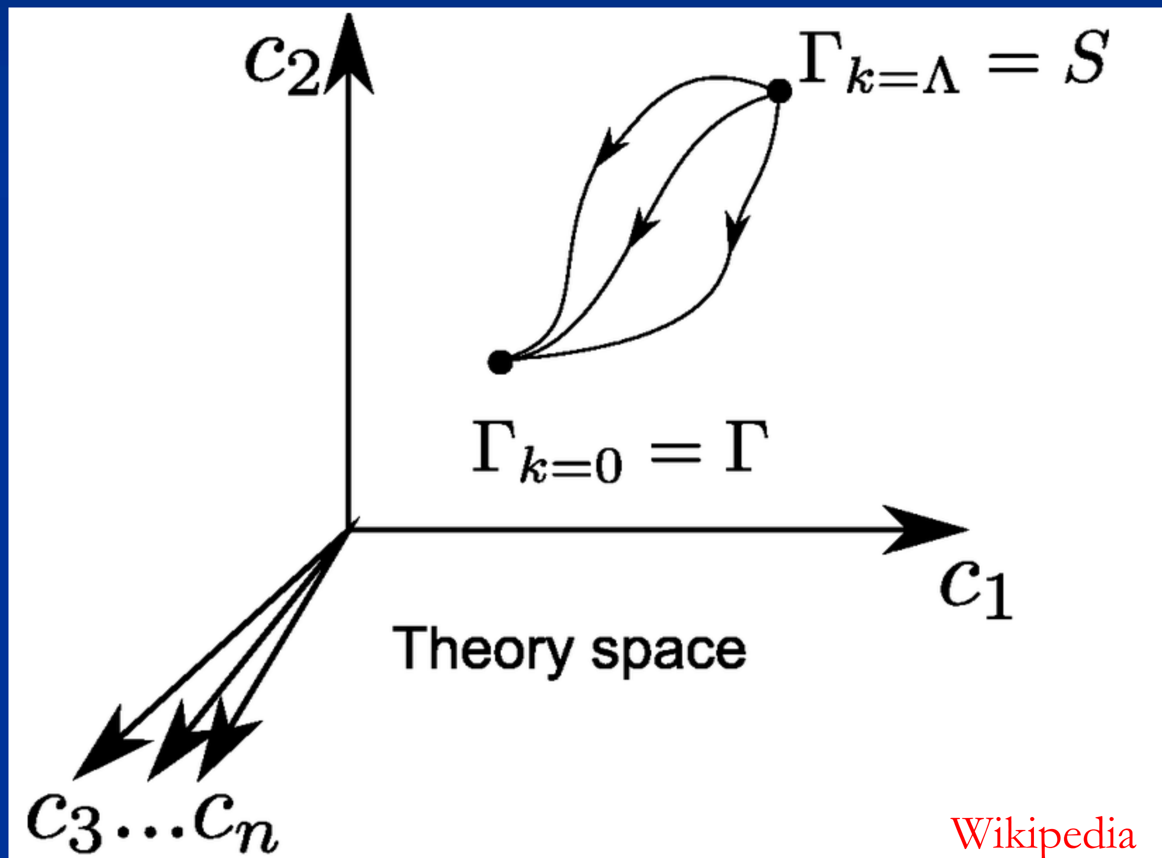




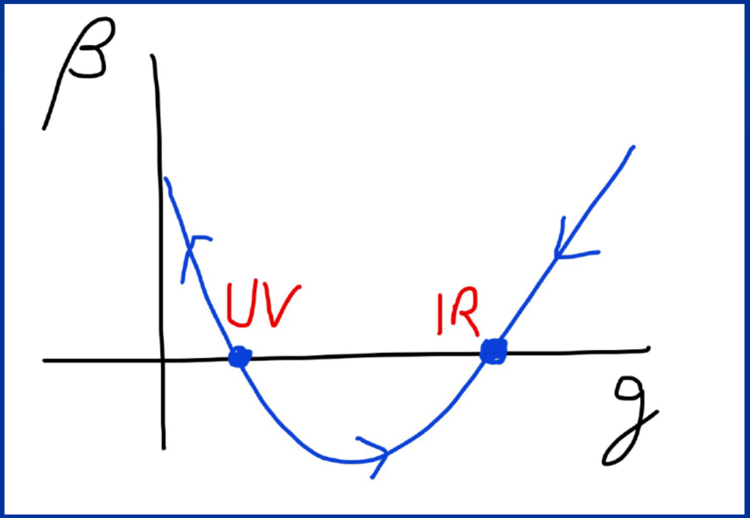
Quantum scale symmetry

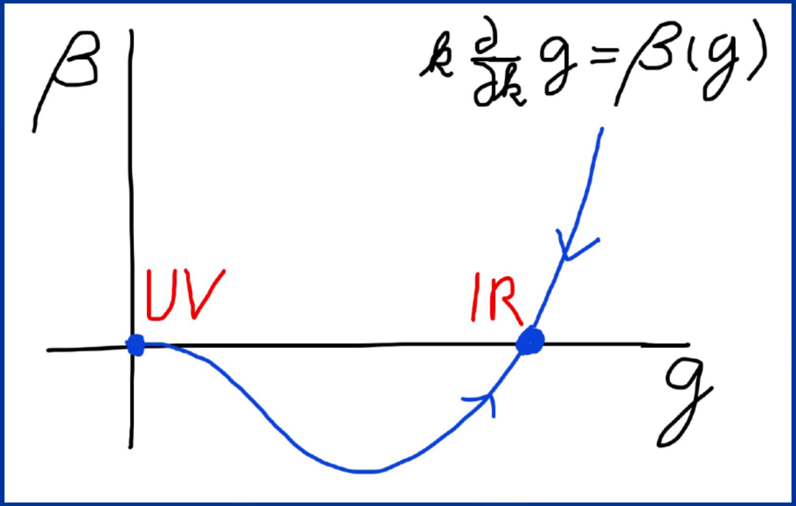
- quantum fluctuations violate scale symmetry
- running dimensionless couplings
- at fixed points , scale symmetry is exact !

functional renormalization : flowing action



Asymptotic safety   Asymptotic freedom





Relevant parameters yield undetermined couplings.
Quartic scalar coupling is not relevant and can therefore be predicted.

Key question (2)

- Why does scale symmetry play a role even though it is violated by quantum fluctuations ?

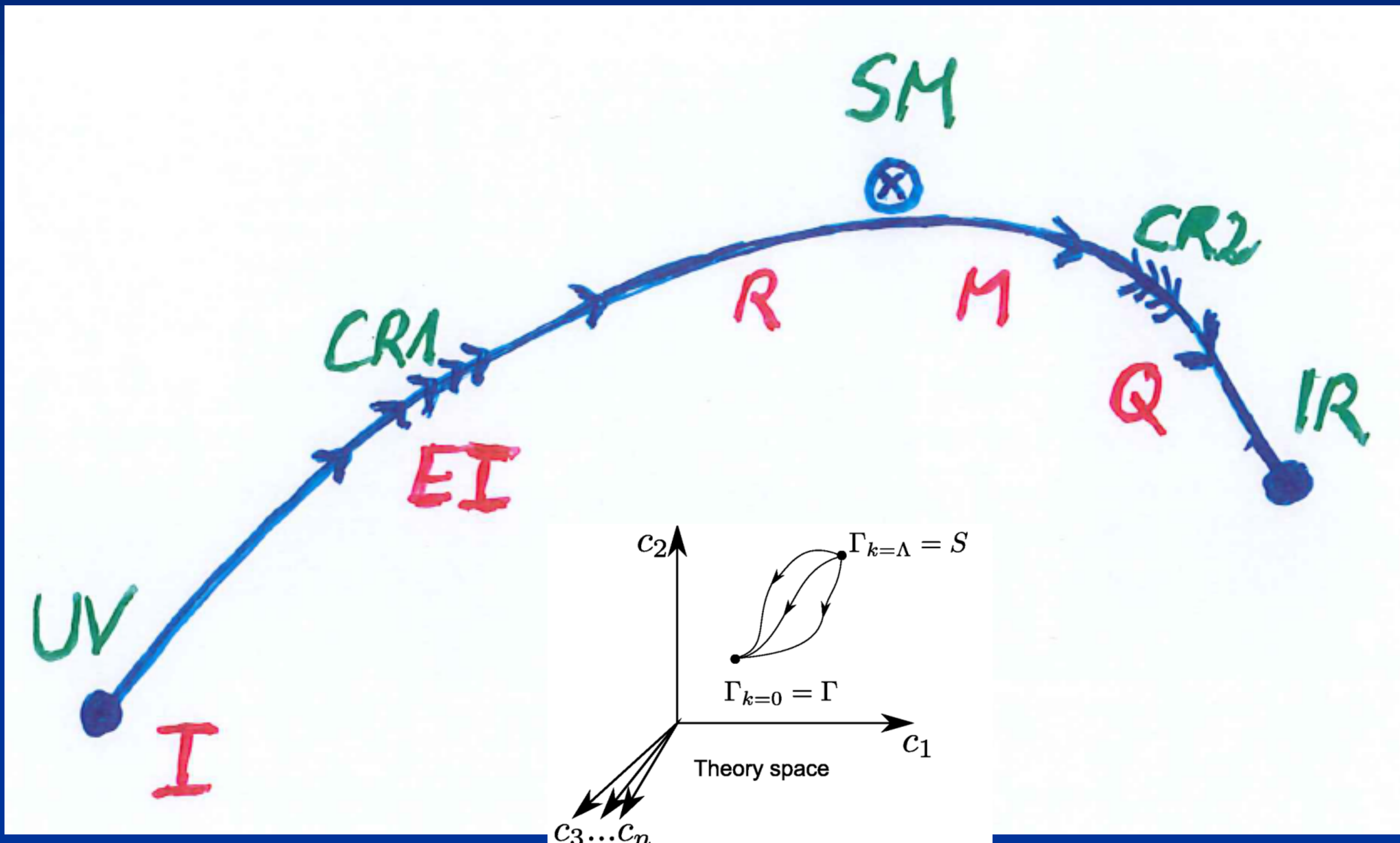
answer

Fixed points
always realize scale symmetry

Key question (3)

- Do fixed points play a role for cosmology ?

Crossover in quantum gravity



Approximate scale symmetry near fixed points

- UV : approximate scale invariance of primordial fluctuation spectrum from inflation

- IR : cosmon is pseudo Goldstone boson of
  spontaneously broken scale symmetry,
  tiny mass,
  responsible for dynamical Dark Energy

Asymptotic safety

if UV fixed point exists :

quantum gravity is

non-perturbatively renormalizable !

S. Weinberg , M. Reuter

a prediction…

Asymptotic safety of gravity and the Higgs boson mass

Mikhail Shaposhnikov

Institut de Théorie des Phénomènes Physiques, École Polytechnique Fédérale de Lausanne, CH-1015 Lausanne, Switzerland

Christof Wetterich

Institut für Theoretische Physik, Universität Heidelberg, Philosophenweg 16, D-69120 Heidelberg, Germany

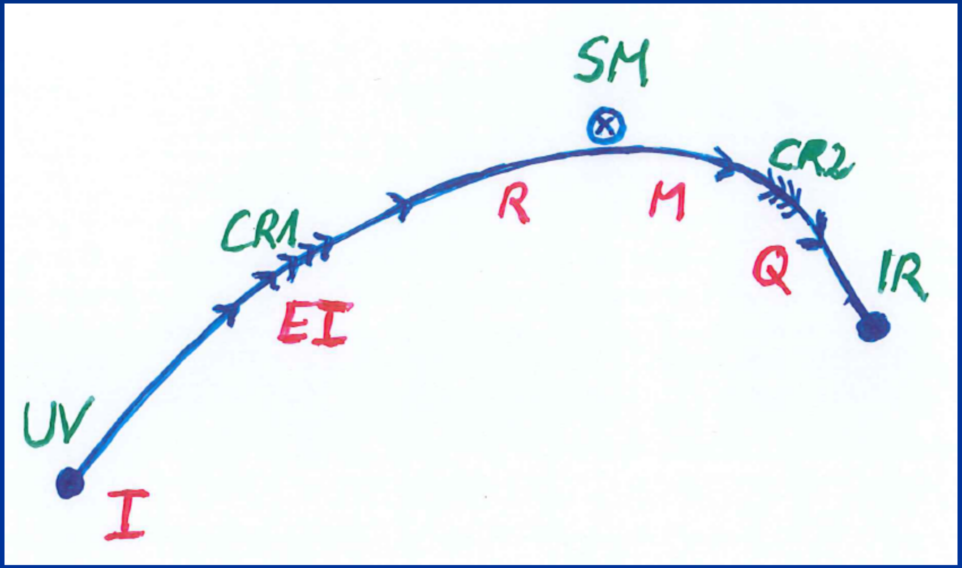
12 January 2010

Abstract

There are indications that gravity is asymptotically safe. The Standard Model (SM) plus gravity could be valid up to arbitrarily high energies. Supposing that this is indeed the case and assuming that there are no intermediate energy scales between the Fermi and Planck scales we address the question of whether the mass of the Higgs boson m_H can be predicted. For a positive gravity induced anomalous dimension $A_\lambda > 0$ the running of the quartic scalar self interaction λ at scales beyond the Planck mass is determined by a fixed point at zero. This results in $m_H = m_{\rm min} = 126$ GeV, with only a few GeV uncertainty. This prediction is independent of the details of the short distance running and holds for a wide class of extensions of the SM as well.

s in $m_H = m_{\rm min} = 126$ GeV, with o

Possible consequences of crossover in quantum gravity



Realistic model for inflation and dark energy with single scalar field

Variable Gravity

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2}\chi^2 R + \mu^2 \chi^2 + \frac{1}{2}\left(B(\chi/\mu) - 6\right)\partial^\mu \chi \partial_\mu \chi \right\}$$

quantum effective action,
variation yields field equations

Einstein gravity : $\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} M^2 R \right\}$

Variable Gravity

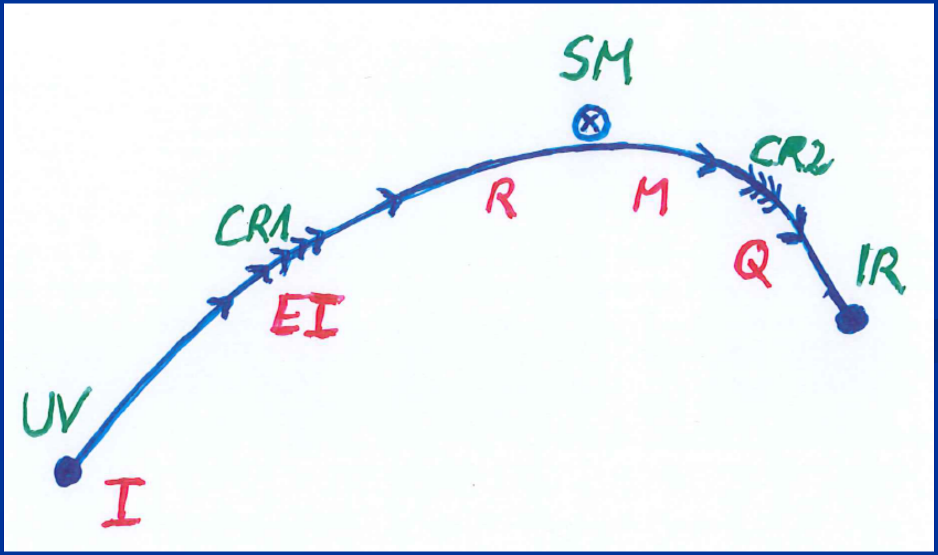
- Scalar field coupled to gravity
- Effective Planck mass depends on scalar field
- Simple quadratic scalar potential involves intrinsic mass μ
- Nucleon and electron mass proportional to dynamical Planck mass
- Neutrino mass has different dependence on scalar field

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2}\chi^2 R + \mu^2\chi^2 + \frac{1}{2}\left(B(\chi/\mu) - 6\right)\partial^\mu\chi\partial_\mu\chi \right\}$$

Cosmological solution : crossover from UV to IR fixed point

- Dimensionless functions as B depend only on ratio μ/χ .
- IR: $\mu \rightarrow 0$, $\chi \rightarrow \infty$
- UV: $\mu \rightarrow \infty$, $\chi \rightarrow 0$

Cosmology makes crossover between fixed points by variation of χ .



renormalization flow and cosmological evolution

- renormalization flow as function of μ

  is mapped by dimensionless functions to

- field dependence of effective action on scalar field χ

  translates by solution of field equation to

- dependence of cosmology an time t or η

Key question (3)

- Do fixed points play a role for cosmology ?

answer

likely

(C) Approximate scale symmetry

Slowly running couplings close to fixed points

$$\frac{\partial g_i}{\partial \ln \chi} = A_i(g_i - g_{i*})$$

Simple mechanism for tiny cosmon-atom couplings

- asymptotic approach to fixed point for dimensionless couplings and mass ratios
- at fixed point : no cosmon coupling to atoms – no time variation of fundamental constants
- very near fixed point : tiny coupling
- how small ?

Neutrino cosmon coupling

- Strong bounds on atom-cosmon coupling from tests of equivalence principle or time variation of couplings.
- No such bounds for neutrino-cosmon coupling.
- In particle physics : Mass generation mechanism for neutrinos differs from charged fermions. Seesaw mechanism involves heavy particles whose mass may depend on the value of the cosmon field.

neutrino mass

$$M_\nu = M_D M_R^{-1} M_D^T + M_L$$

$$M_L = h_L \gamma \frac{d^2}{M_t^2}$$

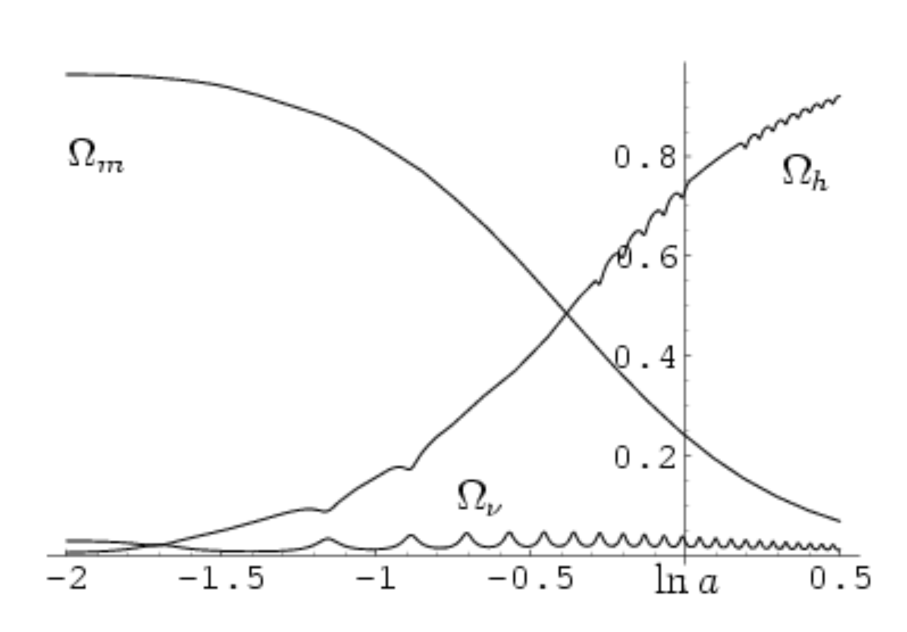
seesaw and cascade mechanism

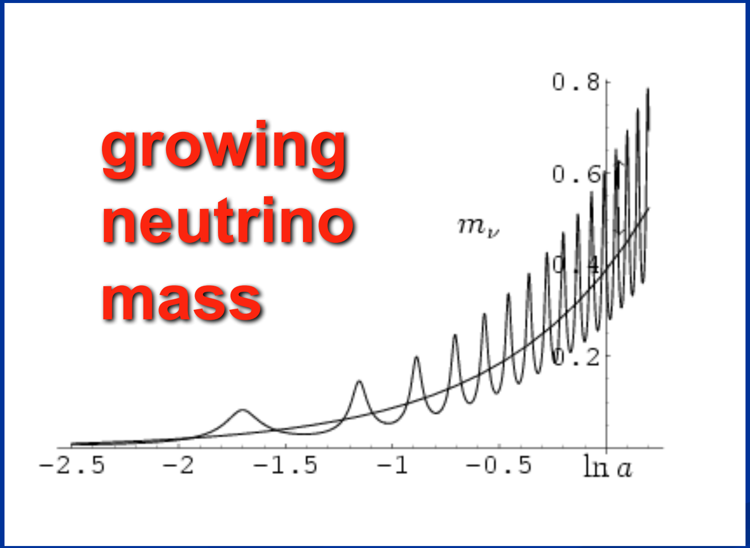
triplet expectation value ~ doublet squared

$$m_\nu = \frac{h_\nu^2 d^2}{m_R} + \frac{h_L \gamma d^2}{M_t^2}$$

omit generation structure

growing neutrino mass triggers transition to almost static dark energy





L.Amendola, M.Baldi,…

connection between dark energy and neutrino properties

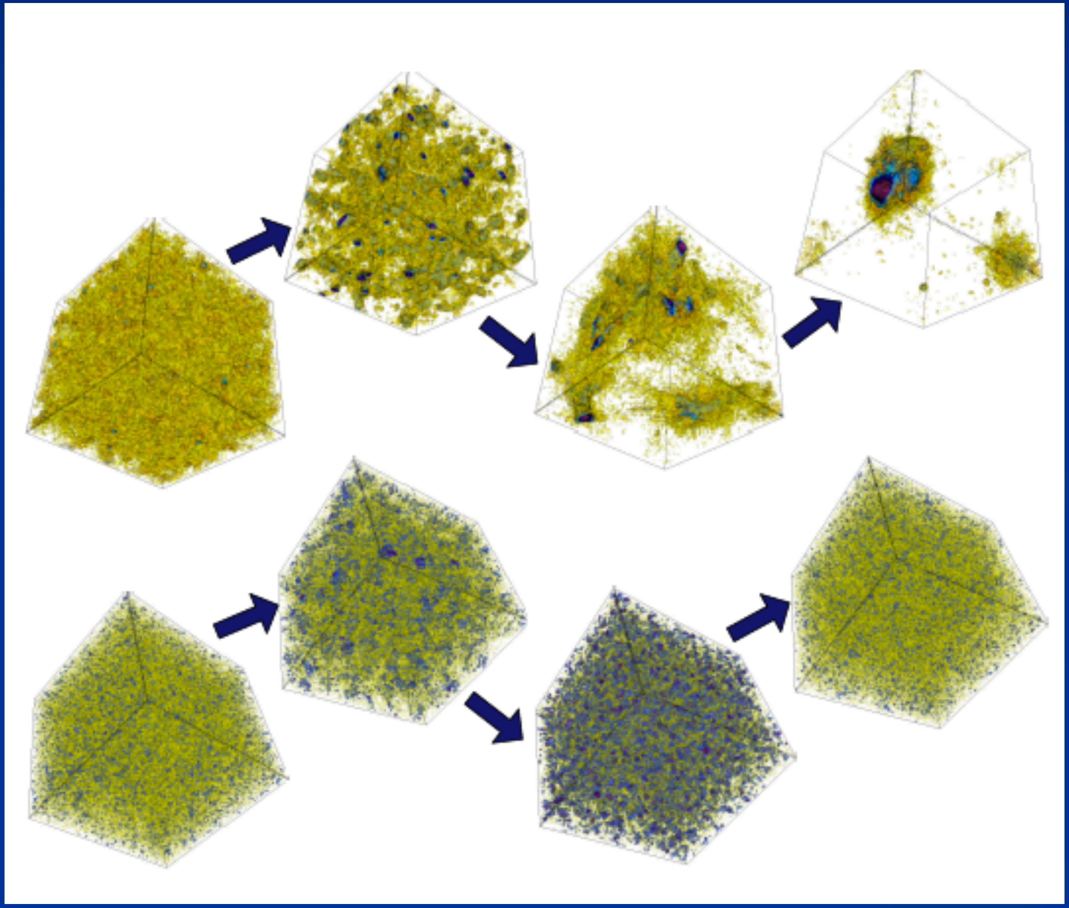
$$[\rho_h(t_0)]^{\frac{1}{4}} = 1.27 \left(\frac{\gamma m_\nu(t_0)}{eV}\right)^{\frac{1}{4}} 10^{-3} eV$$

present dark energy density given by neutrino mass

present equation of state given by neutrino mass !

$$w_0 \approx -1 + \frac{m_\nu(t_0)}{12\text{eV}}$$

Neutrino lumps

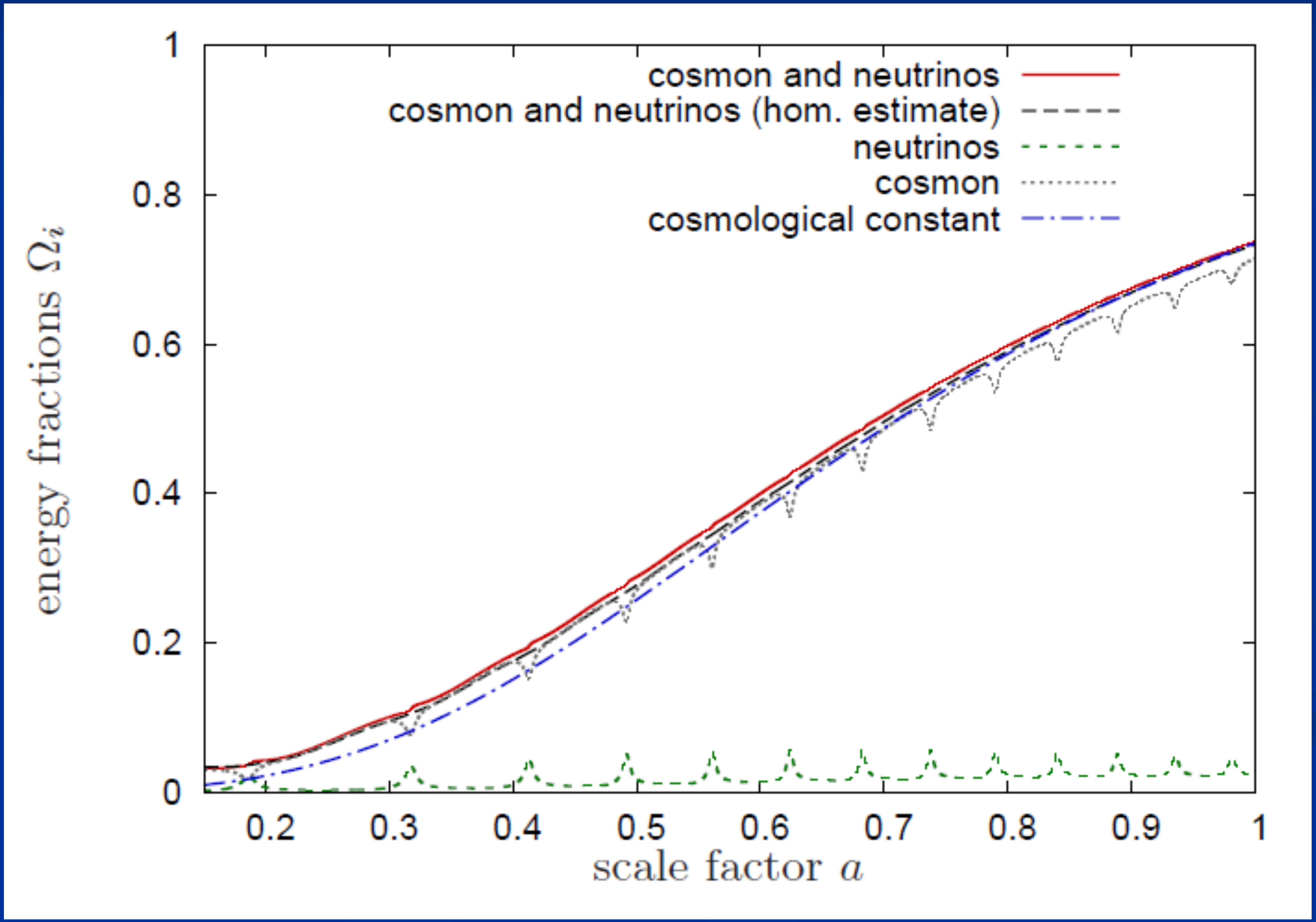


large m_ν

small m_ν

Casas, Pettorino,…

Evolution of dark energy similar to ΛCDM



Crossover cosmology

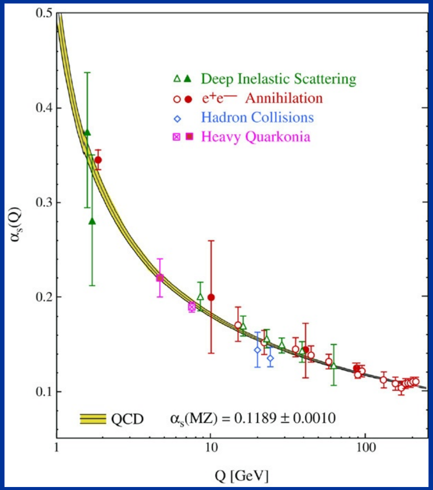
$$\Gamma = \int_x \sqrt{g}\left\{-\frac{1}{2}\chi^2 R + \mu^2\chi^2 + \frac{1}{2}\big(B(\chi/\mu) - 6\big)\partial^\mu\chi\partial_\mu\chi\right\}$$

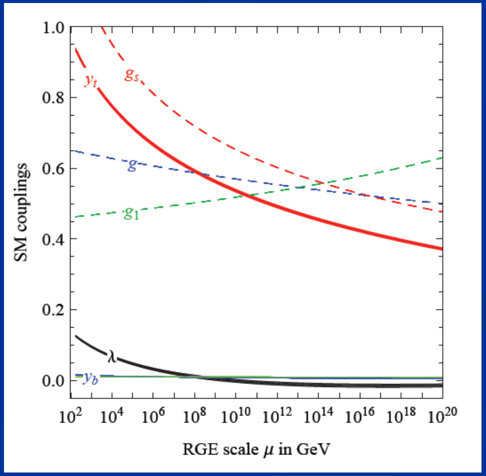
Running coupling B

- B varies if intrinsic scale μ changes

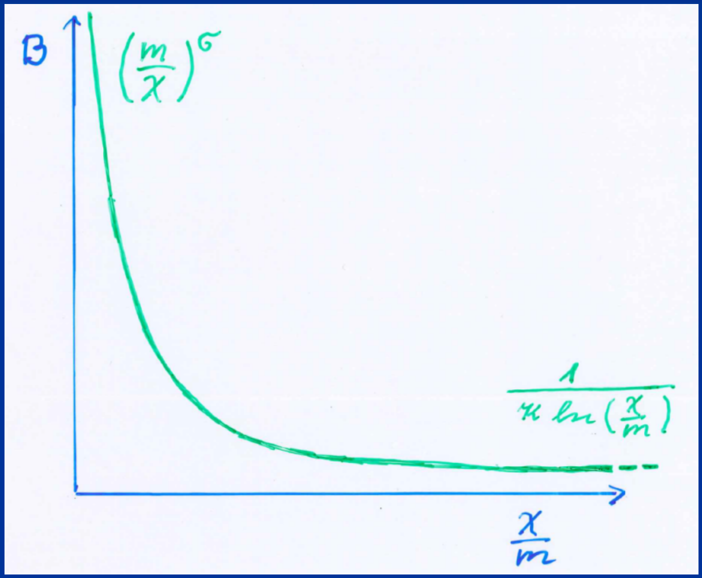
$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2}\chi^2 R + \mu^2 \chi^2 + \frac{1}{2}\left(B(\chi/\mu) - 6\right)\partial^\mu \chi \partial_\mu \chi \right\}$$

- similar to QCD or standard model





Kinetial B :
Crossover between two fixed points



assumption:
running
coupling obeys
flow equation

$$\mu \frac{\partial B}{\partial \mu} = \frac{\kappa \sigma B^2}{\sigma + \kappa B}$$

$$B^{-1} - \frac{\kappa}{\sigma} \ln B = \kappa \left[\ln\left(\frac{\chi}{\mu}\right) - c_t \right] = \kappa \ln\left(\frac{\chi}{m}\right)$$

m : scale of crossover

can be exponentially larger than intrinsic scale μ

Four-parameter model

- model has four dimensionless parameters
- three in kinetial :

  $\sigma \sim 2.5$

  $\varkappa \sim 0.5$

  $c_t \sim 14$ (or m/μ)
- one parameter for growth rate of neutrino mass over electron mass : $\gamma \sim 8$
- + standard model particles and dark matter : sufficient for realistic cosmology from inflation to dark energy
- no more free parameters than ΛCDM

Cosmological solution

- scalar field χ vanishes in the infinite past
- scalar field χ diverges in the infinite future

No tiny dimensionless parameters
(except gauge hierarchy)

- one mass scale $\mu = 2 \cdot 10^{-33}$ eV
- one time scale $\mu^{-1} = 10^{10}$ yr
- Planck mass does not appear as parameter
- Planck mass grows large dynamically

Slow Universe

Asymptotic solution in freeze frame :

$$H = \frac{\mu}{\sqrt{3}} \;,\; \chi = \frac{3^{\frac{1}{4}} m}{2\sqrt{\mu}} (t_c - t)^{-\frac{1}{2}}$$

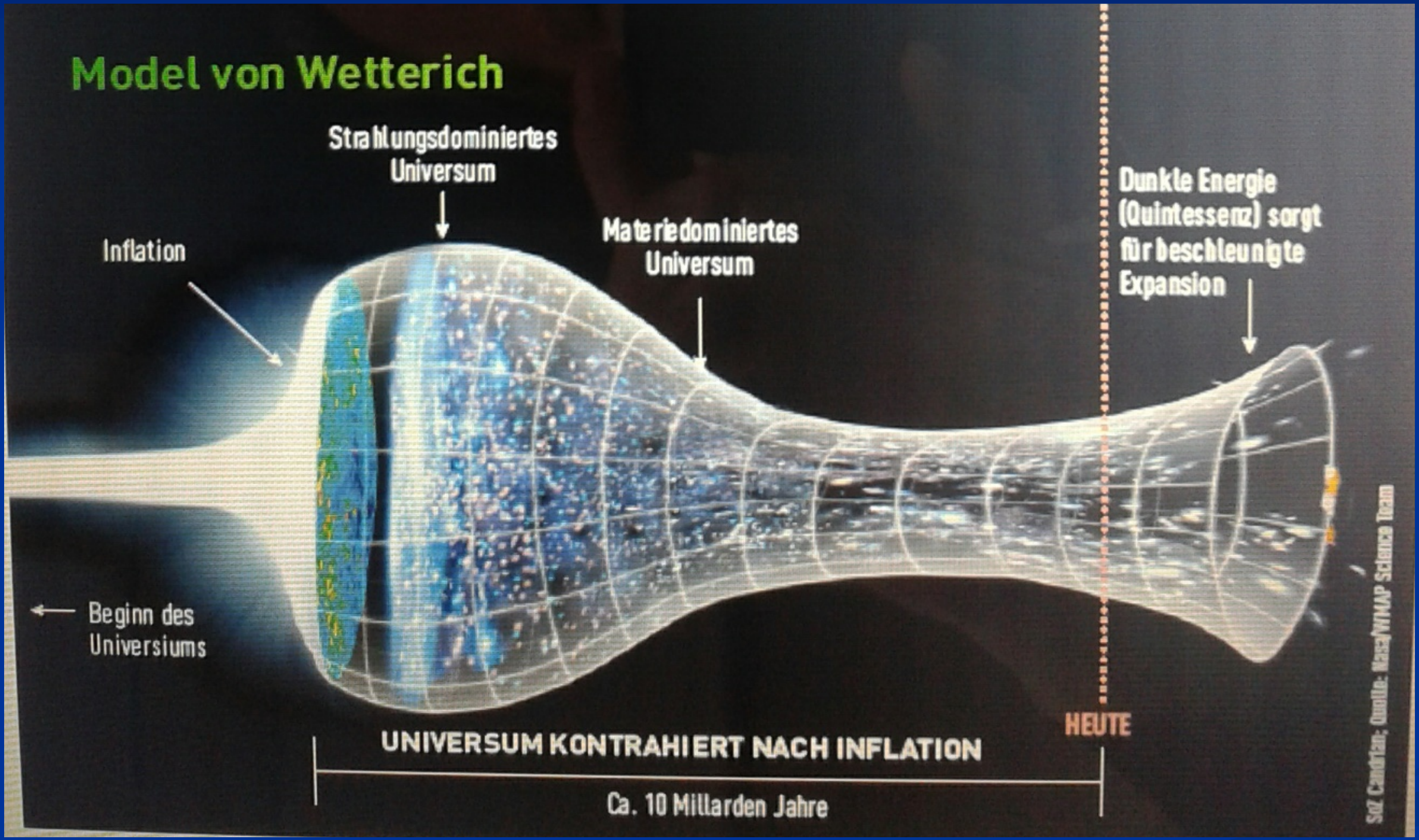
$$\mu = 2 \cdot 10^{-33}\ \mathrm{eV}$$

Expansion or shrinking always slow ,
characteristic time scale of the order of the age of the Universe : $t_{ch} \sim \mu^{-1} \sim$ 10 billion years !

Hubble parameter of the order of present Hubble parameter for all times , including inflation and big bang !

Slow increase of particle masses !

Strange evolution of Universe



Sonntagszeitung Zürich , Laukenmann

Model is compatible with present observations

Together with variation of neutrino mass over electron mass in present cosmological epoch :

model is compatible with all present observations, including inflation and dark energy

$$\Gamma = \int_x \sqrt{g}\left\{-\frac{1}{2}\chi^2 R + \mu^2\chi^2 + \frac{1}{2}\left(B(\chi/\mu) - 6\right)\partial^\mu\chi\partial_\mu\chi\right\}$$

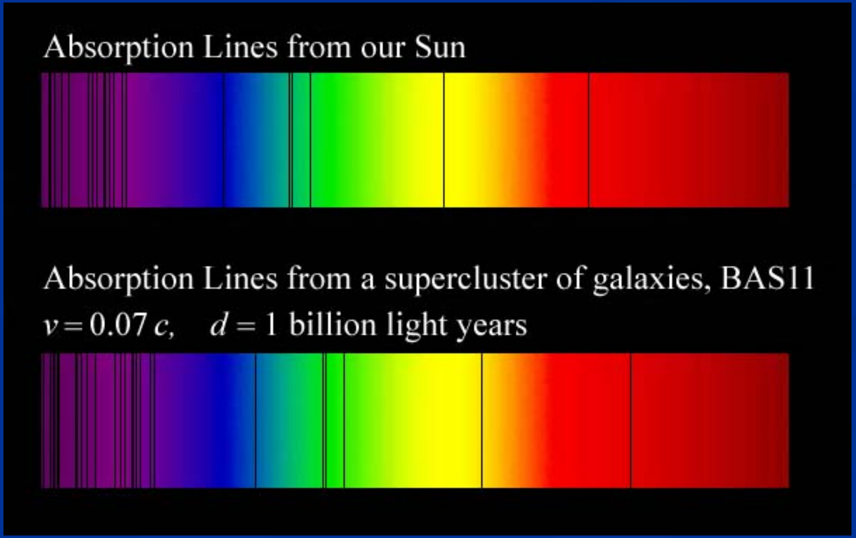
$$B^{-1} - \frac{\kappa}{\sigma}\ln B = \kappa\left[\ln\left(\frac{\chi}{\mu}\right) - c_t\right] = \kappa\ln\left(\frac{\chi}{m}\right)$$

Big bang or freeze ?

Do we know that the Universe expands ?

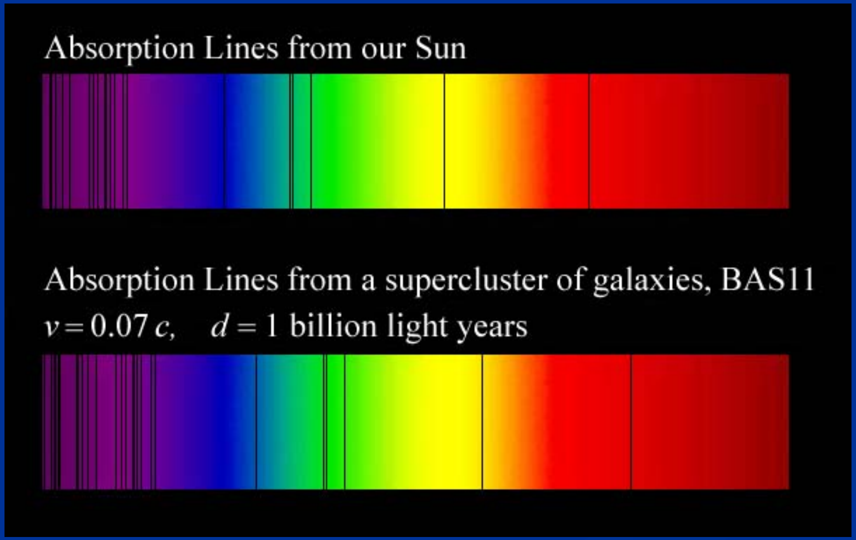
instead of redshift due to expansion :

smaller frequencies have been emitted in the past, because electron mass was smaller !



Why do we see redshift of photons emitted in the distant past ?

photons are more red because they have been emitted with longer wavelength



frequency ~ mass

wavelength ~ atomsize

What is increasing ?

Ratio of distance between galaxies
over size of atoms !

atom size constant : expanding geometry

alternative : shrinking size of atoms

How can particle masses change with time ?

- Particle masses are proportional to scalar field χ . Similar to Higgs field.
- Scalar field varies with time.
- Ratios of particle masses are independent of χ and therefore remain constant.
- Compatibility with observational bounds on time dependence of particle mass ratios.
- Dimensionless couplings are independent of χ .

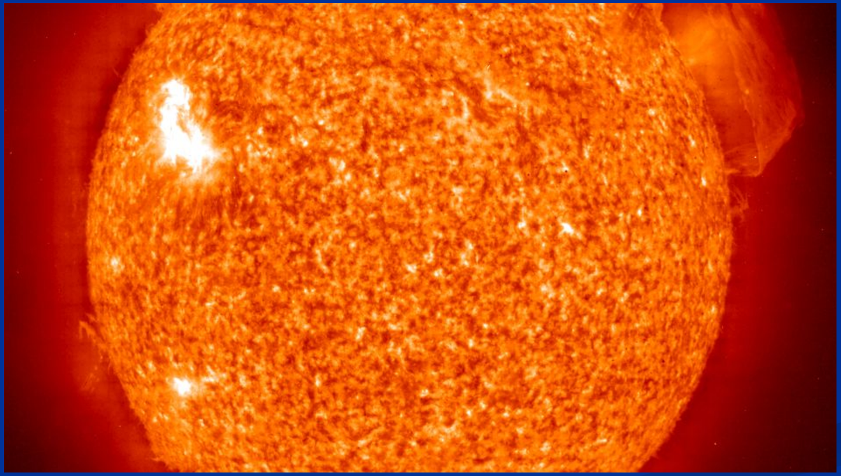
Do we know that the temperature was higher in the early Universe than now ?

Cosmic microwave radiation , nucleosynthesis

instead of

higher temperature :

smaller particle masses

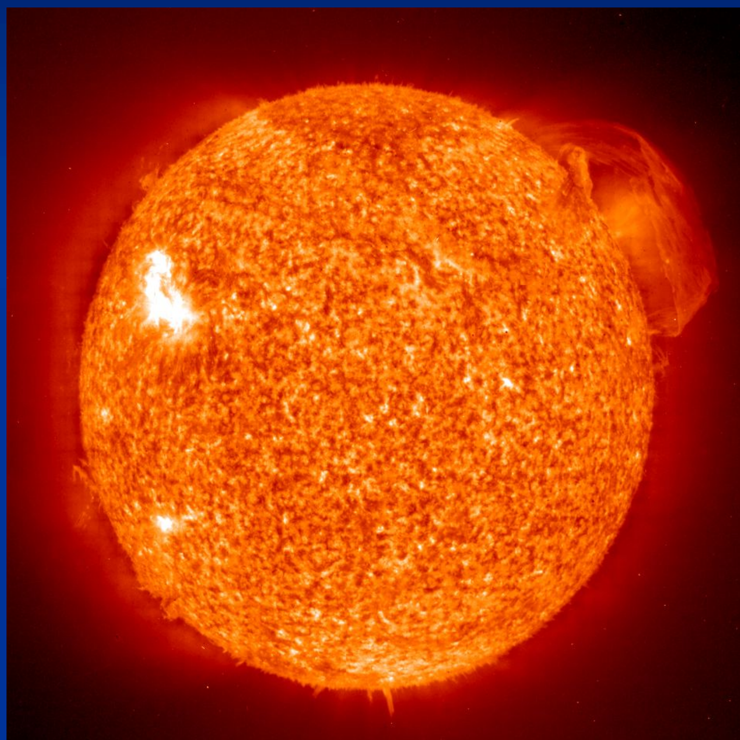


Hot plasma ?

- Temperature in radiation dominated Universe : $T \sim \chi^{1/2}$ smaller than today
- Ratio temperature / particle mass : $T / m_p \sim \chi^{-1/2}$ larger than today
- T/m_p counts ! This ratio decreases with time.

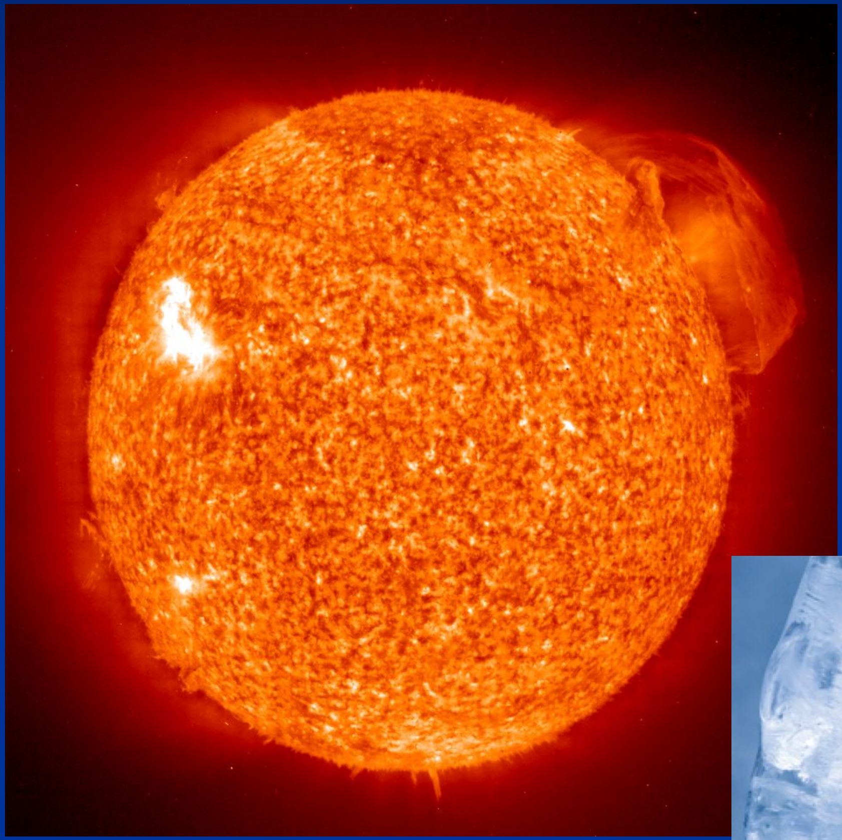
- Nucleosynthesis , CMB emission as in standard cosmology !

Big bang or freeze ?



freeze picture :
only rods for measurements
are set differently !

Big bang or freeze ?



Big bang is not wrong,

but alternative pictures exist !

Conclusions

- Scalar force in cosmology can arise from quantum scale symmetry
- Realistic cosmology

Look for

- time variation of fundamental constants, violation of equivalence principle,
- huge lumps in cosmic neutrino background,
- dynamical dark energy

end

Einstein frame

- "Weyl scaling" maps variable gravity model to Universe with fixed masses and standard expansion history.
- Exact equivalence of different frames !

- Standard gravity coupled to scalar field.

- Only neutrino masses are growing.

Einstein frame

Weyl scaling :

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu} \;,\; \varphi = \frac{2M}{\alpha} \ln\left(\frac{\chi}{\mu}\right)$$

effective action in Einstein frame :

$$\Gamma = \int_x \sqrt{g'} \left\{ -\frac{1}{2} M^2 R' + V'(\varphi) + \frac{1}{2} k^2(\varphi) \partial^\mu \varphi \partial_\mu \varphi \right\}$$

$$V'(\varphi) = M^4 \exp\left(-\frac{\alpha\varphi}{M}\right)$$

$$k^2 = \frac{\alpha^2 B}{4}$$

Field relativity

Weyl scaling :

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu}$$

changes geometry,
not a coordinate transformation

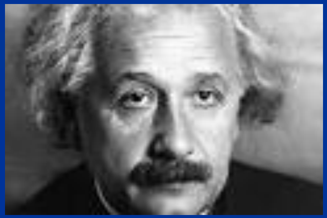
Field relativity : different pictures of cosmology

- same physical content can be described by different pictures
- related by field – redefinitions , e.g. Weyl scaling , conformal scaling of metric
- which picture is usefull ?

Relativity of geometry

- Euclid … Newton : space and time are absolute

- Special relativity : space and time depend on observer
- General relativity : spacetime is influenced by matter (including radiation)



  geometry is independent of coordinates

  geometry is observable
- Field relativity : geometry is relative

Spacetime is a description
of correlations between "matter".

Different pictures exist.

Solution of cosmological constant problem

$$\mathcal{L} = -\frac{1}{2}\chi^2 R + \frac{B-6}{2}\partial^\mu\chi\partial_\mu\chi + \lambda\chi^4$$

not allowed
by quantum gravity

$$\lambda\chi^4 \to (\mu^2\chi^2, \mu^4, \mu^A\chi^{4-A})$$

A ≥ 2

asymptotically vanishing cosmological „constant“

- What matters : Ratio of potential divided by fourth power of Planck mass

$$\frac{V}{\chi^4} = \frac{\mu^2\chi^2}{\chi^4} = \frac{\mu^2}{\chi^2}$$

$V = \mu^2 \chi^2$

- vanishes for $\chi \to \infty$!

small dimensionless number ?

- needs two intrinsic mass scales
- standard approach :V and M (cosmological constant and Planck mass)
- variable gravity : Planck mass moving to infinity , with fixed V ➡ ratio vanishes asymptotically !

no small parameter for dark energy

Einstein frame

Weyl scaling :

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu} \ , \ \varphi = \frac{2M}{\alpha} \ln\left(\frac{\chi}{\mu}\right)$$

effective action in Einstein frame :

$$\Gamma = \int_x \sqrt{g'} \left\{ -\frac{1}{2} M^2 R' + V'(\varphi) + \frac{1}{2} k^2(\varphi) \partial^\mu \varphi \partial_\mu \varphi \right\}$$

$$V'(\varphi) = M^4 \exp\left(-\frac{\alpha\varphi}{M}\right)$$

$$k^2 = \frac{\alpha^2 B}{4}$$

Quintessence

Dynamical dark energy ,

generated by scalar field (cosmon)

C.Wetterich,Nucl.Phys.B302(1988)668, 24.9.87
P.J.E.Peebles,B.Ratra,ApJ.Lett.325(1988)L17, 20.10.87

Prediction :

homogeneous dark energy influences recent cosmology

- of same order as dark matter -

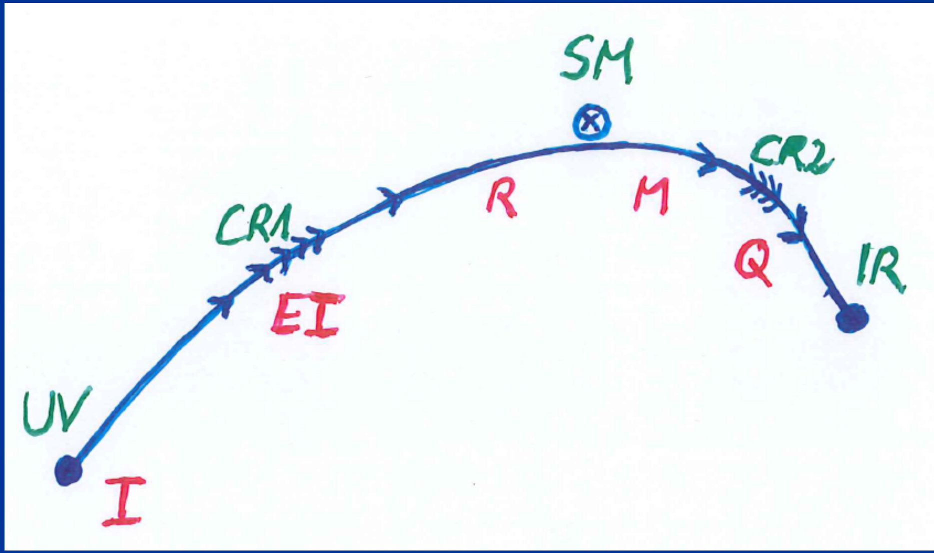
Original models do not fit the present observations

…. modifications

(different growth of neutrino mass)

Second stage of crossover

- from SM to IR
- in sector Beyond Standard Model
- affects neutrino masses first (seesaw or cascade mechanism)



Varying particle masses at onset of second crossover

- All particle masses **except for neutrinos** are proportional to χ .
- Ratios of particle masses remain constant.
- Compatibility with observational bounds on time dependence of particle mass ratios.
- Neutrino masses show stronger increase with χ , such that **ratio neutrino mass over electron mass grows** .

Cosmic trigger

- Stop of evolution of scalar field when neutrinos become non-relativistic
- Transition from scaling solution to (almost) cosmological constant

neutrino mass

$$M_\nu = M_D M_R^{-1} M_D^T + M_L$$

(?) ....
C.Wetterich, Nucl.Phys.B187 (1981) 343

$$M_L = h_L \gamma \frac{d^2}{M_t^2}$$

cascade (seesaw II) mechanism

M.Magg, C.W. 1980

Neutrino cosmon coupling

- realized by dependence of neutrino mass on value of cosmon field

$$\beta(\varphi) = -M \frac{\partial}{\partial \varphi} \ln m_\nu(\varphi)$$

- $\beta \approx 1$: cosmon mediated attractive force between neutrinos has similar strength as gravity

growing neutrinos change cosmon evolution

$$\ddot{\varphi} + 3H\dot{\varphi} = -\frac{\partial V}{\partial \varphi} + \frac{\beta(\varphi)}{M}(\rho_\nu - 3p_\nu),$$

$$\beta(\varphi) = -M\frac{\partial}{\partial \varphi}\ln m_\nu(\varphi) = \frac{M}{\varphi - \varphi_t}$$

modification of conservation equation for neutrinos

$$\dot{\rho}_\nu + 3H(\rho_\nu + p_\nu) = -\frac{\beta(\varphi)}{M}(\rho_\nu - 3p_\nu)\dot{\varphi}$$

$$= -\frac{\dot{\varphi}}{\varphi - \varphi_t}(\rho_\nu - 3p_\nu)$$

effective cosmological trigger for stop of cosmon evolution : neutrinos get non-relativistic

- this has happened recently !
- sets scales for dark energy !

In quantum gravity,
the graviton fluctuations can
play an important role on
distances as large as the
size of the Universe

- for long range scalar fields and dynamical dark energy
- not for all quantities

Instability of graviton propagator

effective action $$\Gamma = \int_x \sqrt{g}\left(-\frac{M^2}{2}R + V\right)$$

flat space: $$G^{-1} = \frac{M^2\, q^2}{4} - \frac{V}{2}$$

Instability for V>0 : ”tachyonic mass term” $-\frac{2V}{M^2}$

curved space: $$G^{-1} = \sqrt{g}\left\{\frac{M^2}{4}\left(-D^2 + \frac{2R}{3}\right) - \frac{V}{2}\right\}$$

Graviton barrier

Quantum gravity computation :

For $\chi \to \infty$

V cannot increase stronger than M^2 !

Instability of graviton propagator is avoided

Graviton barrier and solution of the cosmological constant problem

V cannot increase stronger than M^2 !

If M increases with χ, and for cosmological solutions where χ asymptotically diverges for time going to infinity:

Effective cosmological constant vanishes in infinite future

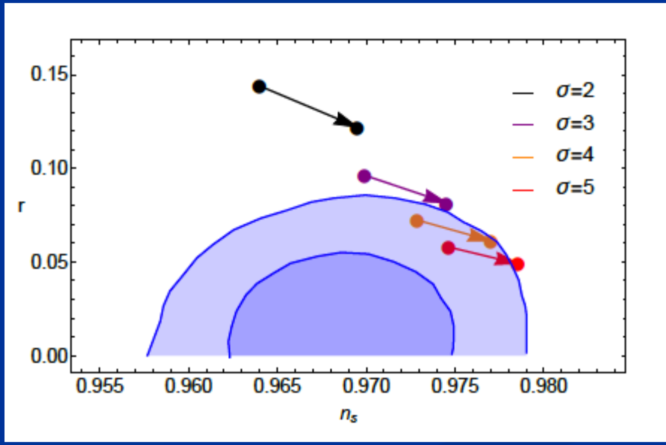
$$\mathbf{M} = \chi \quad : \quad \mathbf{V} = \mu^2 \chi^2$$

Simplicity

simple description of all cosmological epochs

natural incorporation of Dark Energy :

- inflation
- Early Dark Energy
- present Dark Energy dominated epoch



J.Rubio…

Why should you care about the freeze picture of the Universe ?

Some aspects are understood easier :

- Beginning of Universe
- Role of scale symmetry
- Range of impact of quantum gravity

infinite past

Infinite past : slow inflation

$\sigma = 2$: field equations

$$\ddot{\chi} + \left(3H + \frac{1}{2}\frac{\dot{\chi}}{\chi}\right)\dot{\chi} = \frac{2\mu^2\chi^2}{m}$$

$$H = \sqrt{\frac{\mu^2}{3} + \frac{m\dot{\chi}^2}{6\chi^3}} - \frac{\dot{\chi}}{\chi}$$

approximative solution

$$H = \frac{\mu}{\sqrt{3}}\ ,\ \chi = \frac{3^{\frac{1}{4}}m}{2\sqrt{\mu}}(t_c - t)^{-\frac{1}{2}}$$

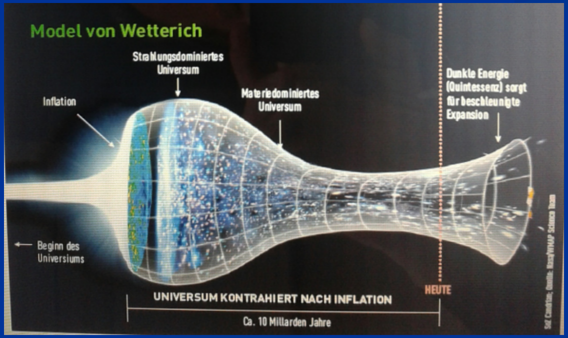
particles become massless in infinite past !

Eternal Universe

Asymptotic solution in freeze frame :

$$H = \frac{\mu}{\sqrt{3}} \; , \;\; \chi = \frac{3^{\frac{1}{4}} m}{2\sqrt{\mu}} (t_c - t)^{-\frac{1}{2}}$$

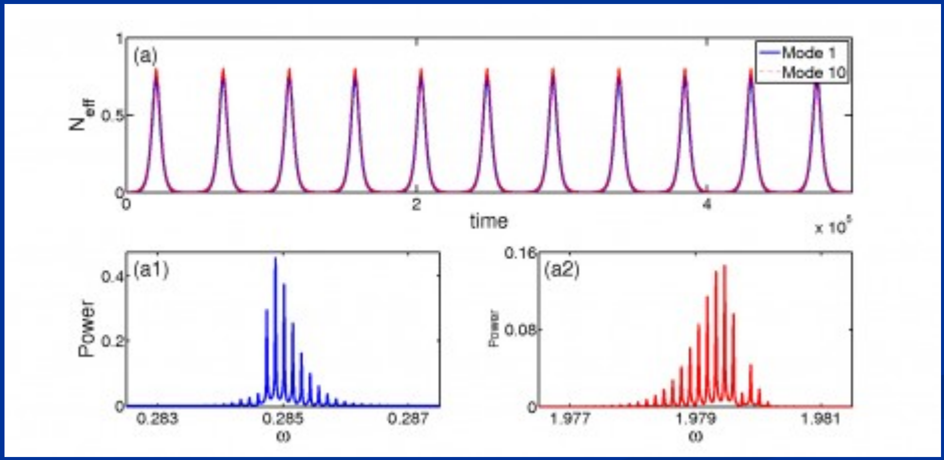
- solution valid back to the infinite past in physical time
- no singularity



- physical time to infinite past is infinite

Physical time

count oscillations ….



Physical time

field equation for scalar field mode

$$(\partial_\eta^2 + 2Ha\partial_\eta + k^2 + a^2m^2)\varphi_k = 0$$

$$\varphi_k = \frac{\tilde{\varphi}_k}{a} \qquad \left\{\partial_\eta^2 + k^2 + a^2\left(m^2 - \frac{R}{6}\right)\right\}\tilde{\varphi}_k = 0$$

determine physical time by counting number of oscillations

$$\tilde{t}_p = n_k \qquad n_k = \frac{k\eta}{\pi}$$

(m=0)

Physical time

- counting : discrete
- invariant under field transformations
- same in all frames

Big bang singularity in Einstein frame is field singularity !

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu} \ , \ \varphi = \frac{2M}{\alpha} \ln\left(\frac{\chi}{\mu}\right)$$

choice of frame with constant particle masses is not well suited if physical masses go to zero !

Cosmon coupling to atoms

has to be much weaker than gravity

$$\frac{m_n}{M} = f_0 + c_n \exp\left(-\beta_n \frac{\varphi}{M}\right)$$

requires for small time variation of couplings : $\left|\frac{c_n}{f_0}\right| \ll 1$ or $|\beta_n| \ll 1$

similar for fine structure constant

$$\mathcal{L}_F = \frac{1}{4}\left[\frac{1}{e_0^2} + c_\alpha \exp\left(-\beta_\alpha \frac{\varphi}{M}\right)\right] F_{\mu\nu}F^{\mu\nu}$$

$$\frac{1}{e^2} = \frac{1}{e_0^2} + c_\alpha \exp\left(-\beta_\alpha \frac{\varphi}{M}\right) = \frac{1}{4\pi\alpha}$$